



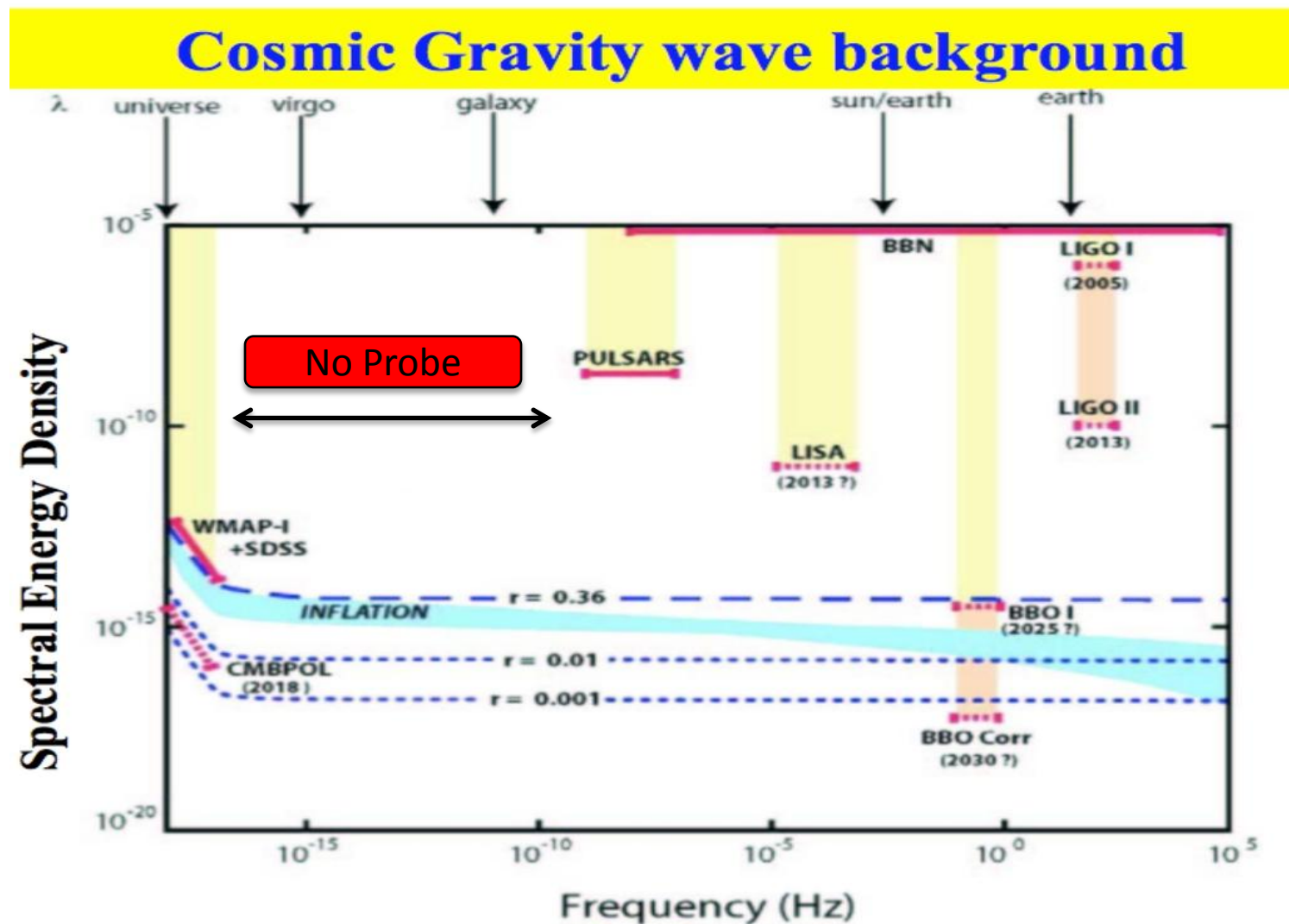
A New Window into Stochastic Gravity Waves

Aditya Rotti

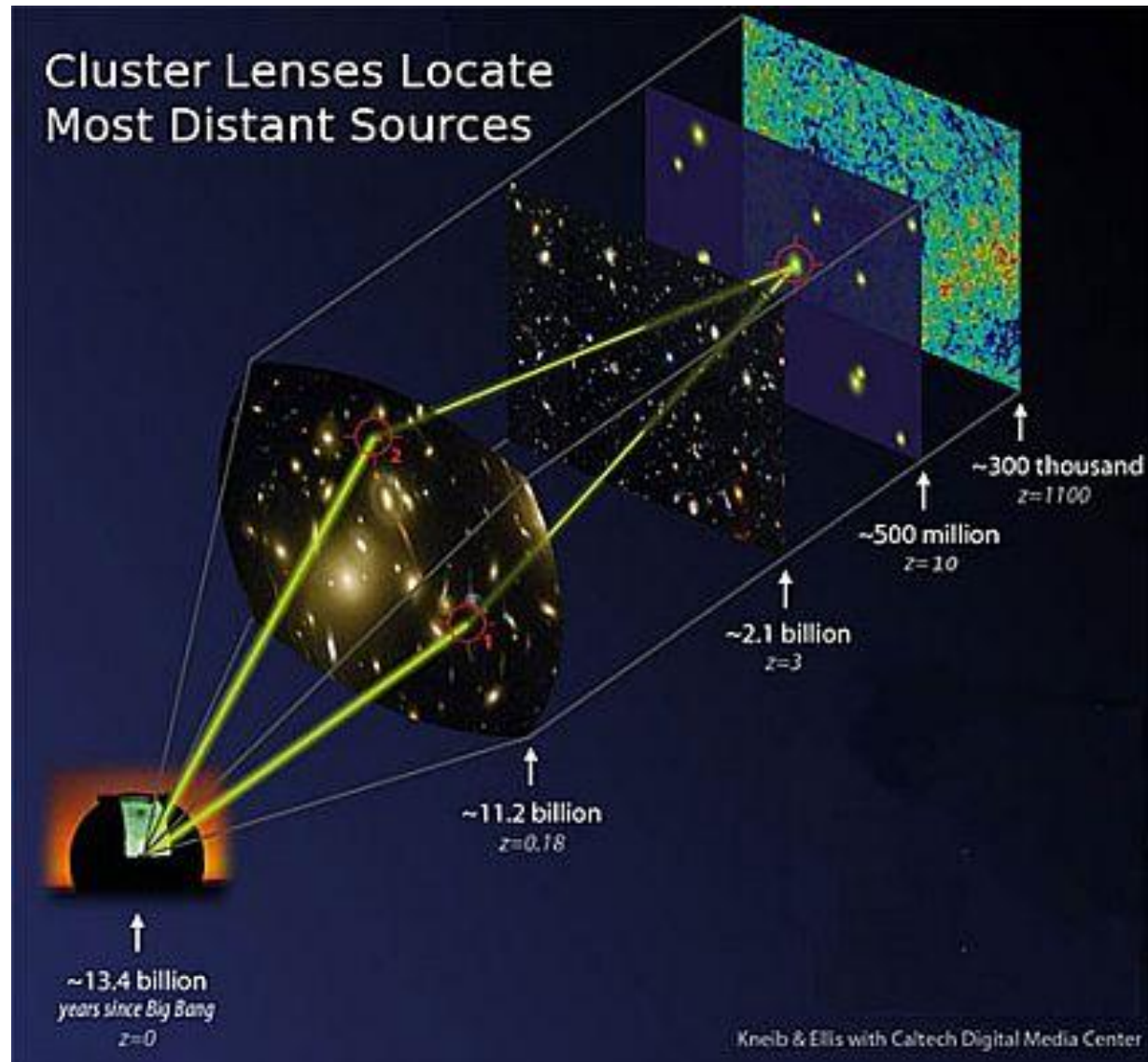
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Indo-UK Meet
11th August 2011

The Current Landscape



Lensing



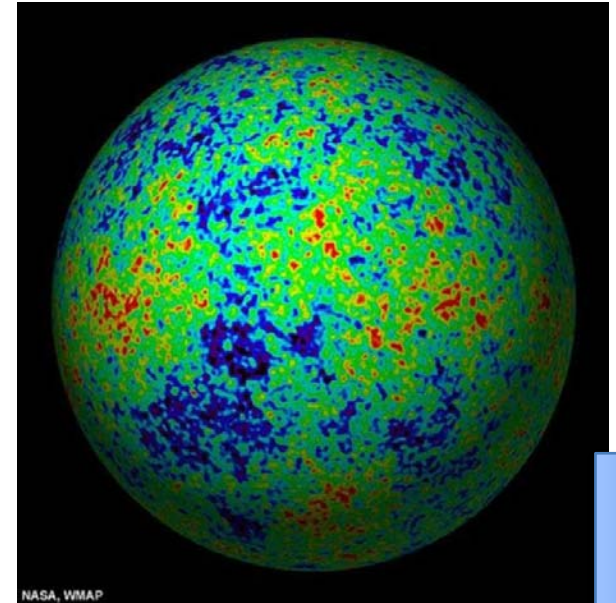
Lensing

Direction of photon arrival changed

$$(\theta_0, \phi_0) \rightarrow (\theta_0 + \delta\theta, \phi_0 + \delta\phi)$$

Let Δ denote the displacement on the sphere.

$$\Delta_a = - \sum_{lm} (h_{lm}^\oplus Y_{lm:a} + h_{lm}^\otimes Y_{lm:b} \epsilon^b_a)$$



Angular power spectrum of photon displacements

$$C_l^{h^\oplus} = \frac{1}{2l+1} \sum_{m=-l}^{m=+l} \langle h_{lm}^\oplus h_{lm}^{\oplus*} \rangle$$

Gradient spectrum

$$C_l^{h^\otimes} = \frac{1}{2l+1} \sum_{m=-l}^{m=+l} \langle h_{lm}^\otimes h_{lm}^{\otimes*} \rangle$$

Curl spectrum

Lensing Modifications to CMB

Lensing induces power transfer between the two polarization spectra

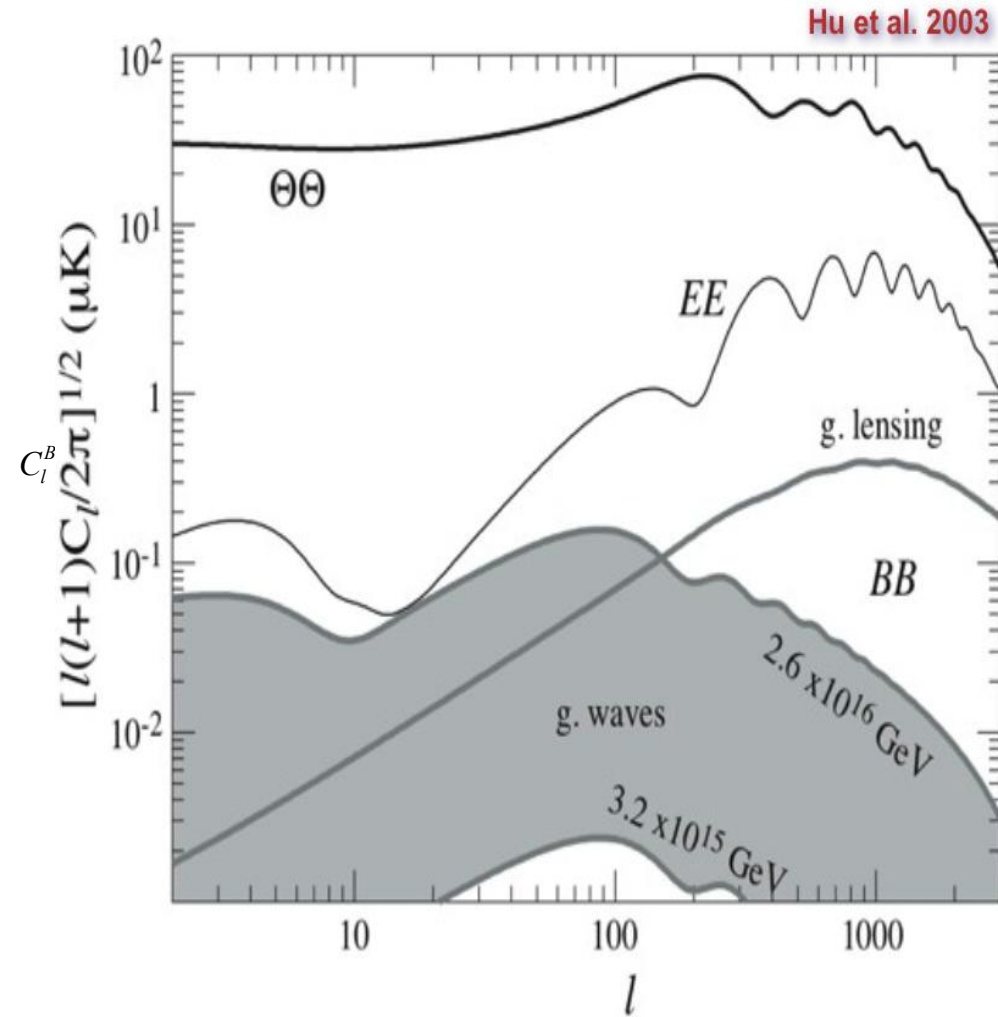
$$\begin{pmatrix} \tilde{C}_l^{EE} \\ \tilde{C}_l^{BB} \end{pmatrix} = \begin{pmatrix} A_{11} & A_{12} \\ A_{21} & A_{22} \end{pmatrix} \begin{pmatrix} C_l^{EE} \\ C_l^{BB} \end{pmatrix}$$

Lensing Kernel

$$\begin{aligned} C_l^{\tilde{E}} &= C_l^E - (l^2 + l - 4)RC_l^E + \frac{1}{2(2l+1)} \\ &\times \sum_{l_1 l_2} \left[C_{l_1}^{h^\oplus} ({}_2F_{ll_1 l_2}^\oplus)^2 + C_{l_1}^{h^\otimes} ({}_2F_{ll_1 l_2}^\otimes)^2 \right] [(C_{l_2}^E + C_{l_2}^B) + (-1)^L (C_{l_2}^E - C_{l_2}^B)] \\ C_l^{\tilde{B}} &= C_l^B - (l^2 + l - 4)RC_l^B + \frac{1}{2(2l+1)} \\ &\times \sum_{l_1 l_2} \left[C_{l_1}^{h^\oplus} ({}_2F_{ll_1 l_2}^\oplus)^2 + C_{l_1}^{h^\otimes} ({}_2F_{ll_1 l_2}^\otimes)^2 \right] [(C_{l_2}^E + C_{l_2}^B) - (-1)^L (C_{l_2}^E - C_{l_2}^B)] \end{aligned}$$

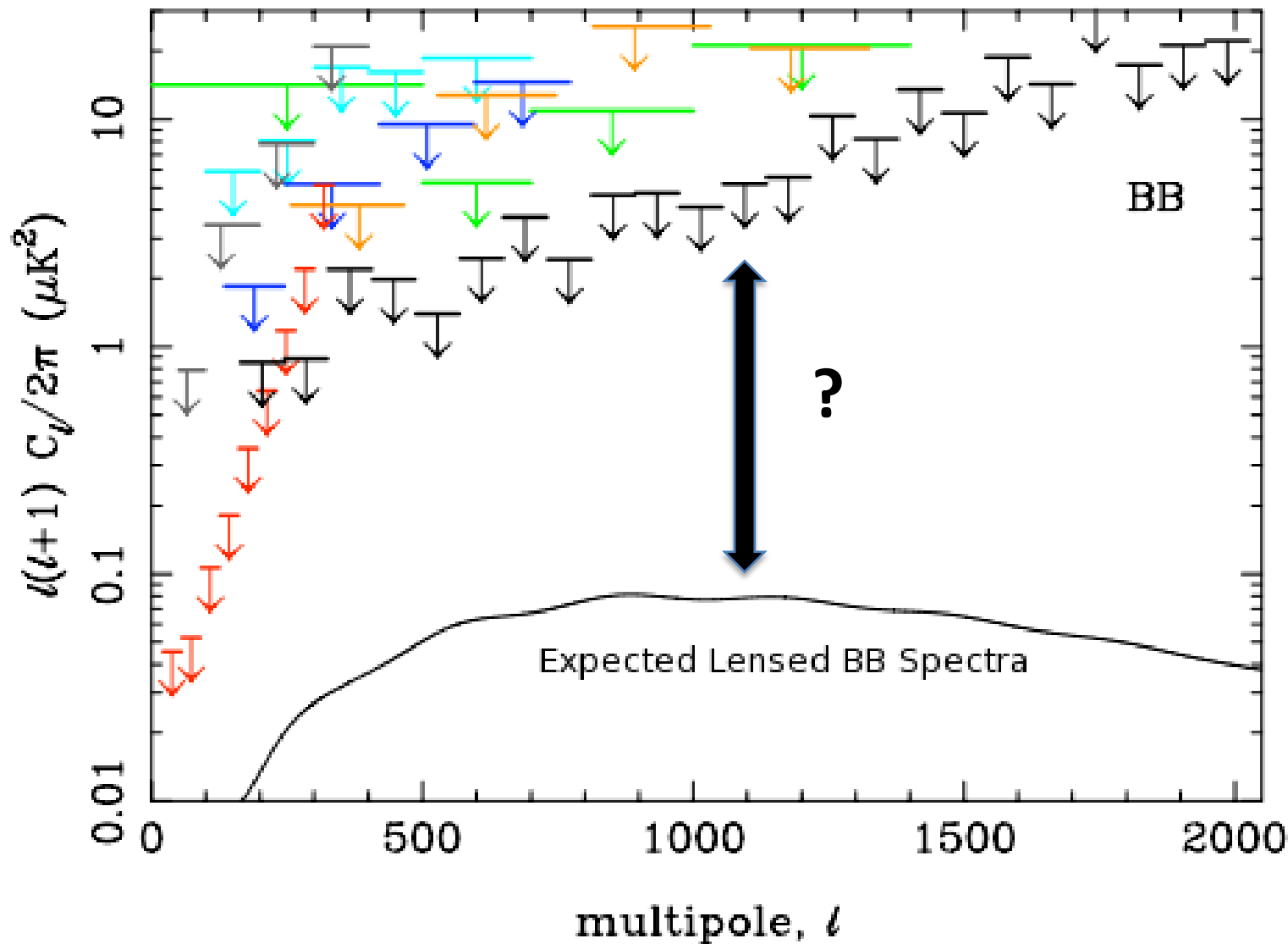
Hints !!

- Current lensing considerations are only due to scalar perturbations.
- The scalar power spectrum is already well measured.
- A huge difference between current upper limits on the BB spectra and the expected signal.

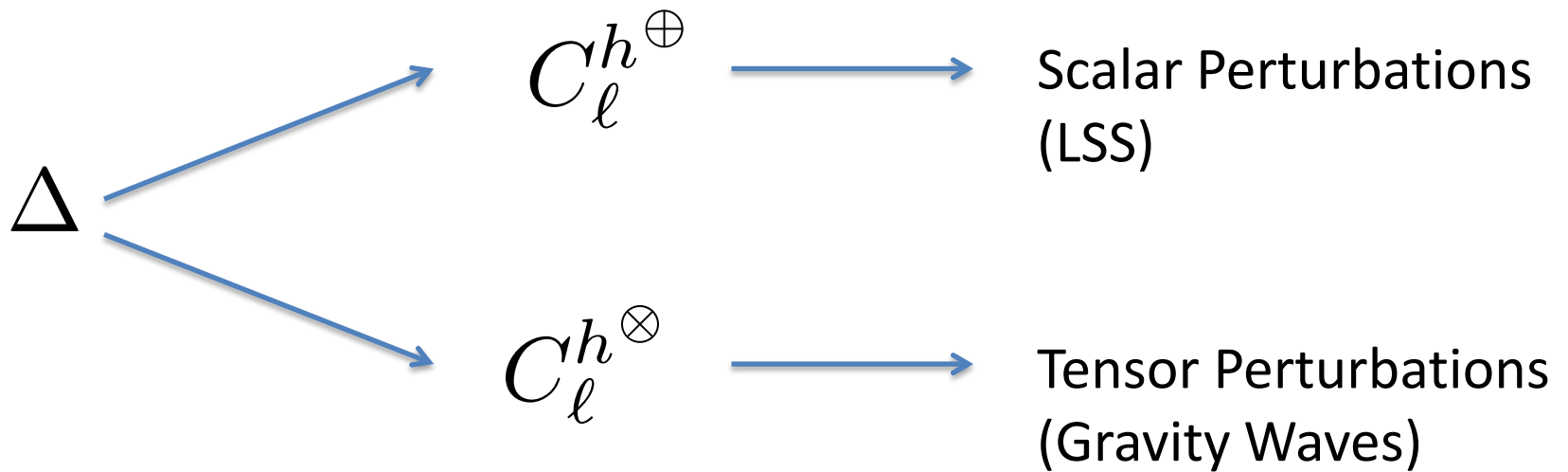


Hints !!

QUAD/BICEP



Just a Reminder



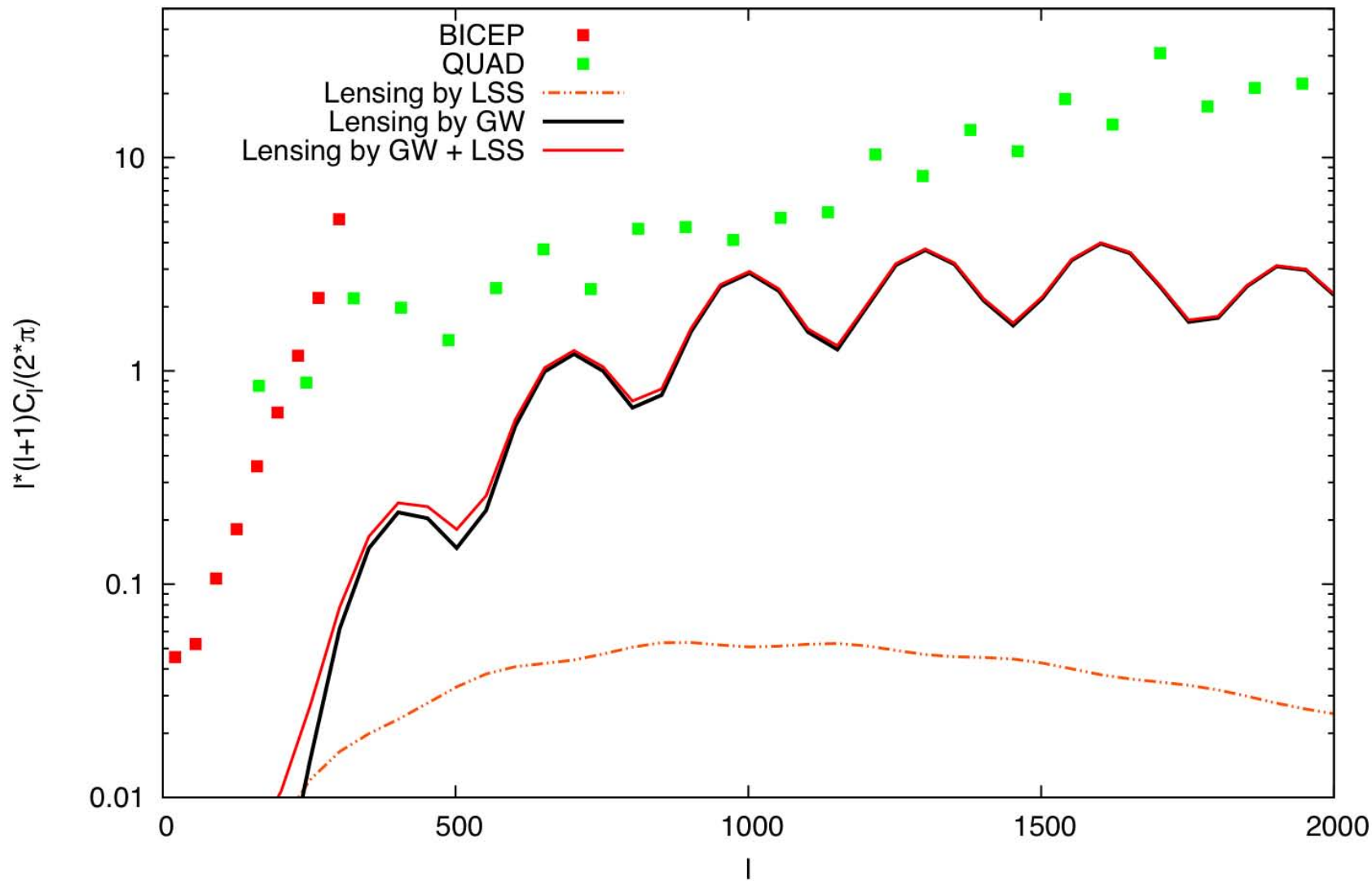
Toy model :

$$C_l^{h^\otimes} = C_l^{h^\oplus}$$

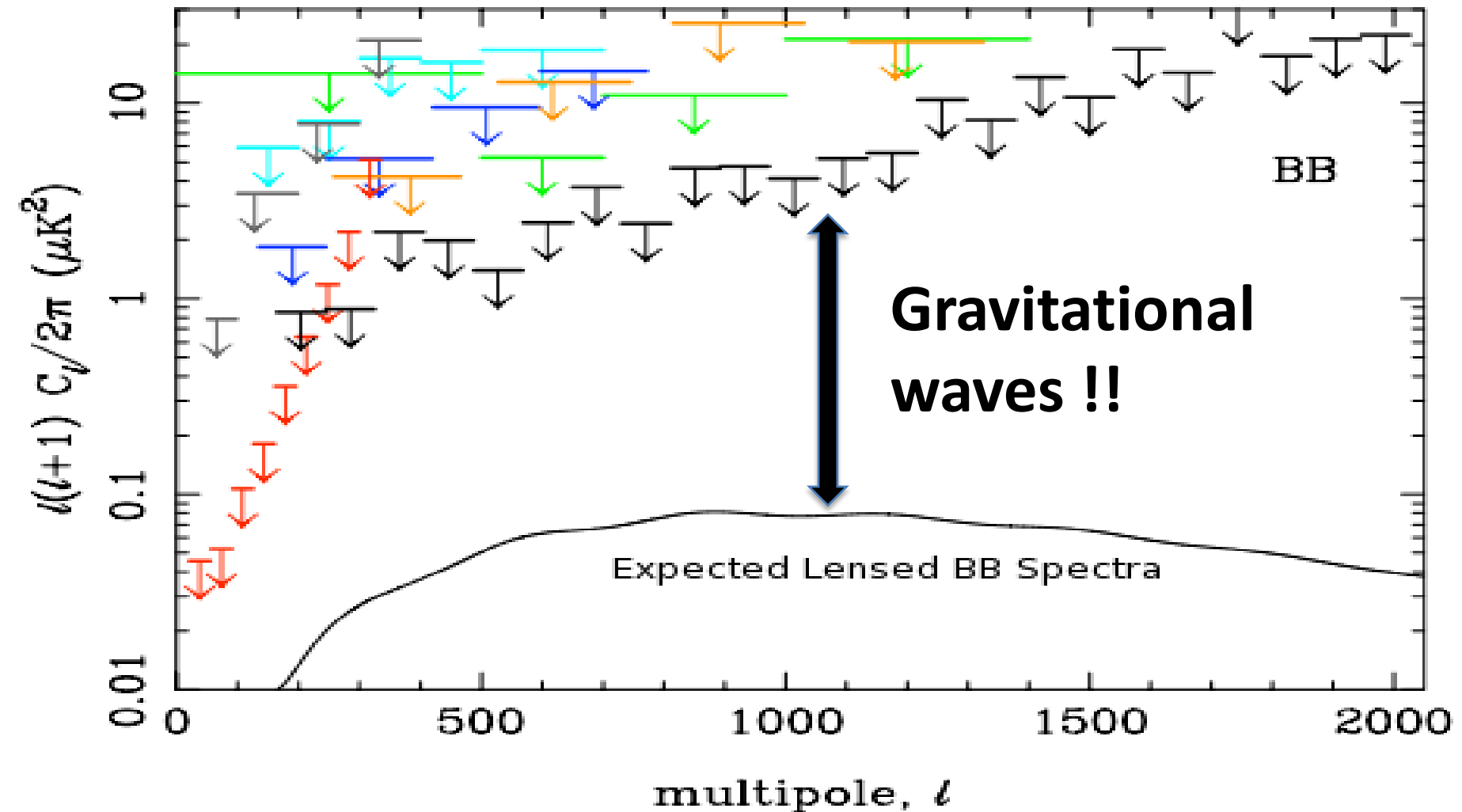
Power Spectrum of the projected lensing potential $C_l^{\psi\psi}$

A Comparison : GW more efficient !

C_l^{BB} from C_l^{EE} due to Lensing by Scalar and Tensor Perturbations



Can we constrain Gravitational Waves Power Spectra ?

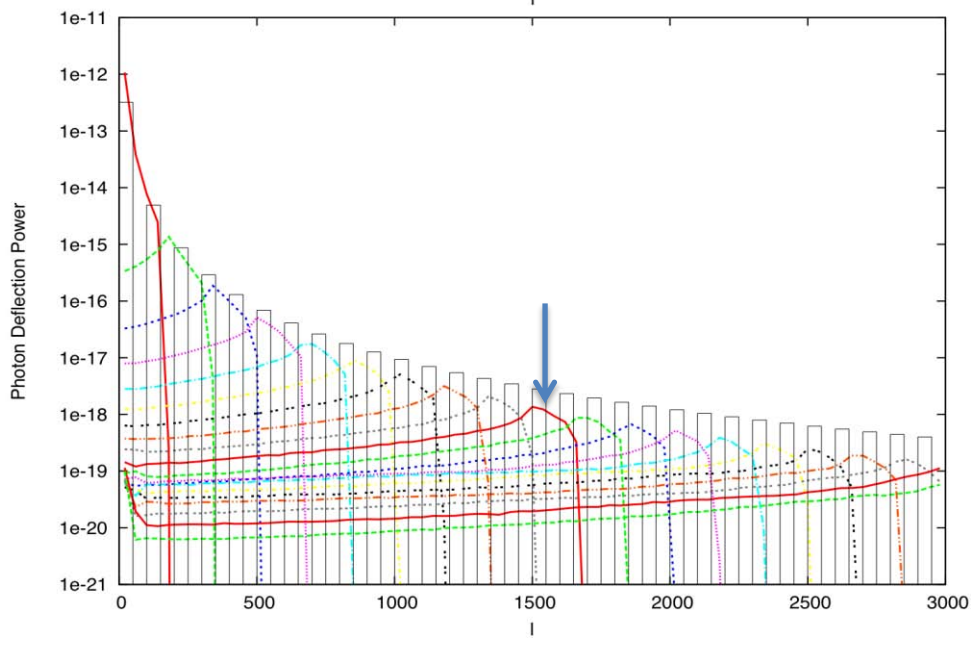
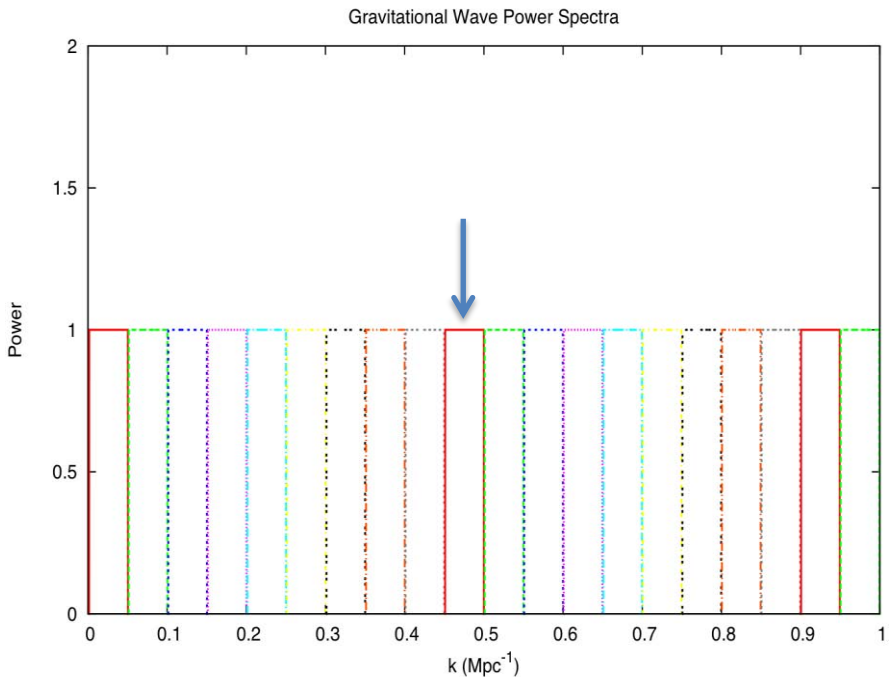
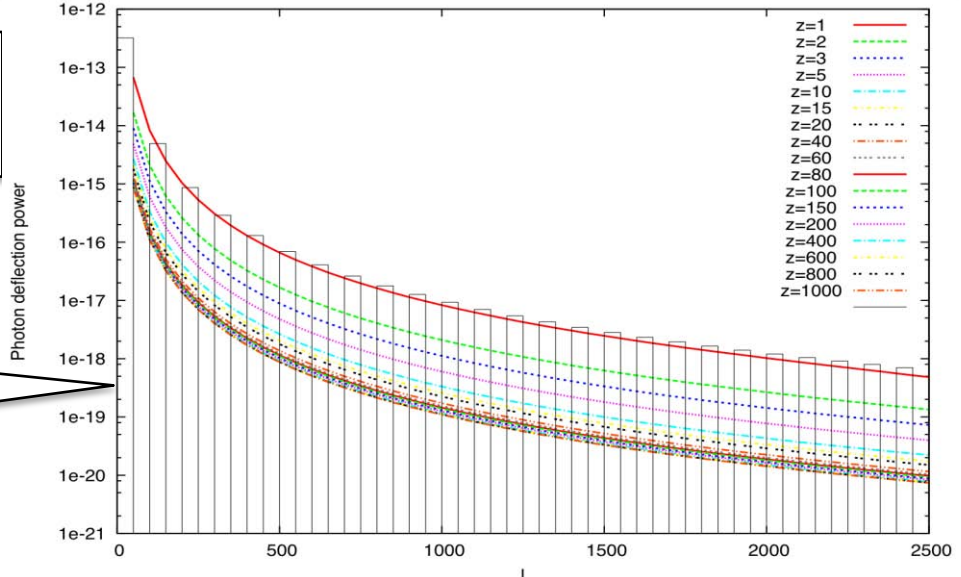


GW Power Spectra ↔ Deflection Power Spectra

$$C_l^{h\otimes} = \frac{\pi}{l^2(l+1)^2} \frac{(l+2)!}{(l-2)!} \int d^3\mathbf{k} P_T(k, z) |T_{eff}|^2$$

C. Li and A. Cooray, PRD
74, 023521(2006)

P(k)=Constant
GW sourced at different z

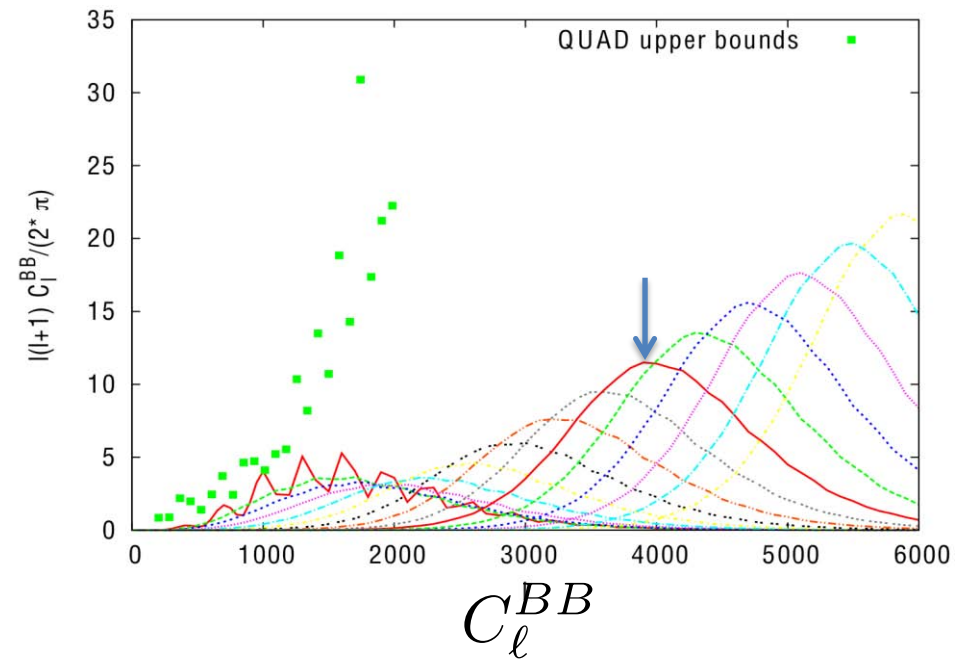
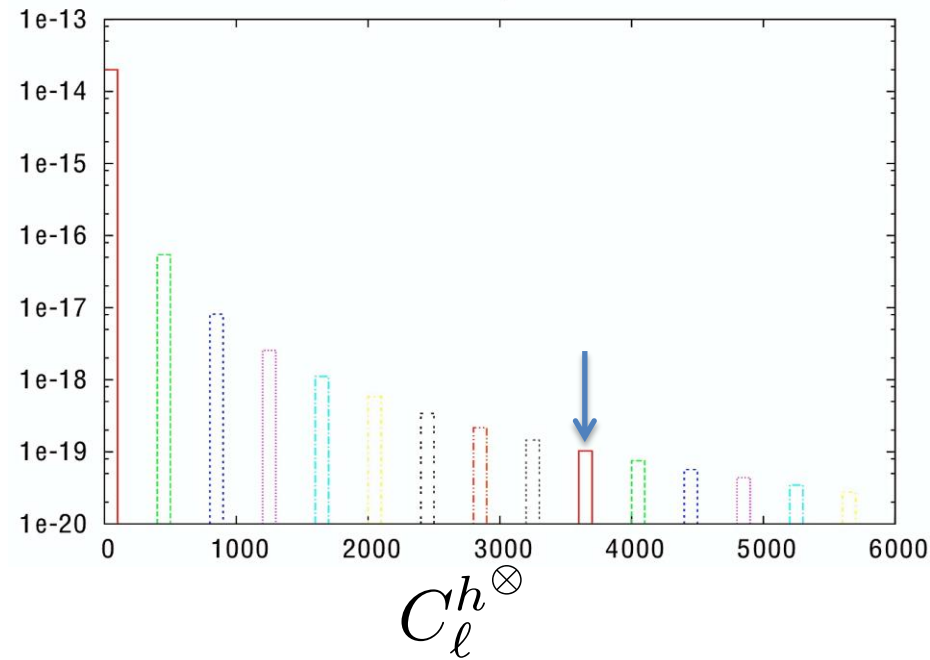


Constraining the deflection power spectra

Sensitivity of the lensed BB spectra to the deflection angular power spectra !!

The **best constraints** on the deflection power in the **specific bins** will come from the measurements of the C_l^{BB} spectra at the location of the **respective peaks**.

rms deflection power --> $1e-6$

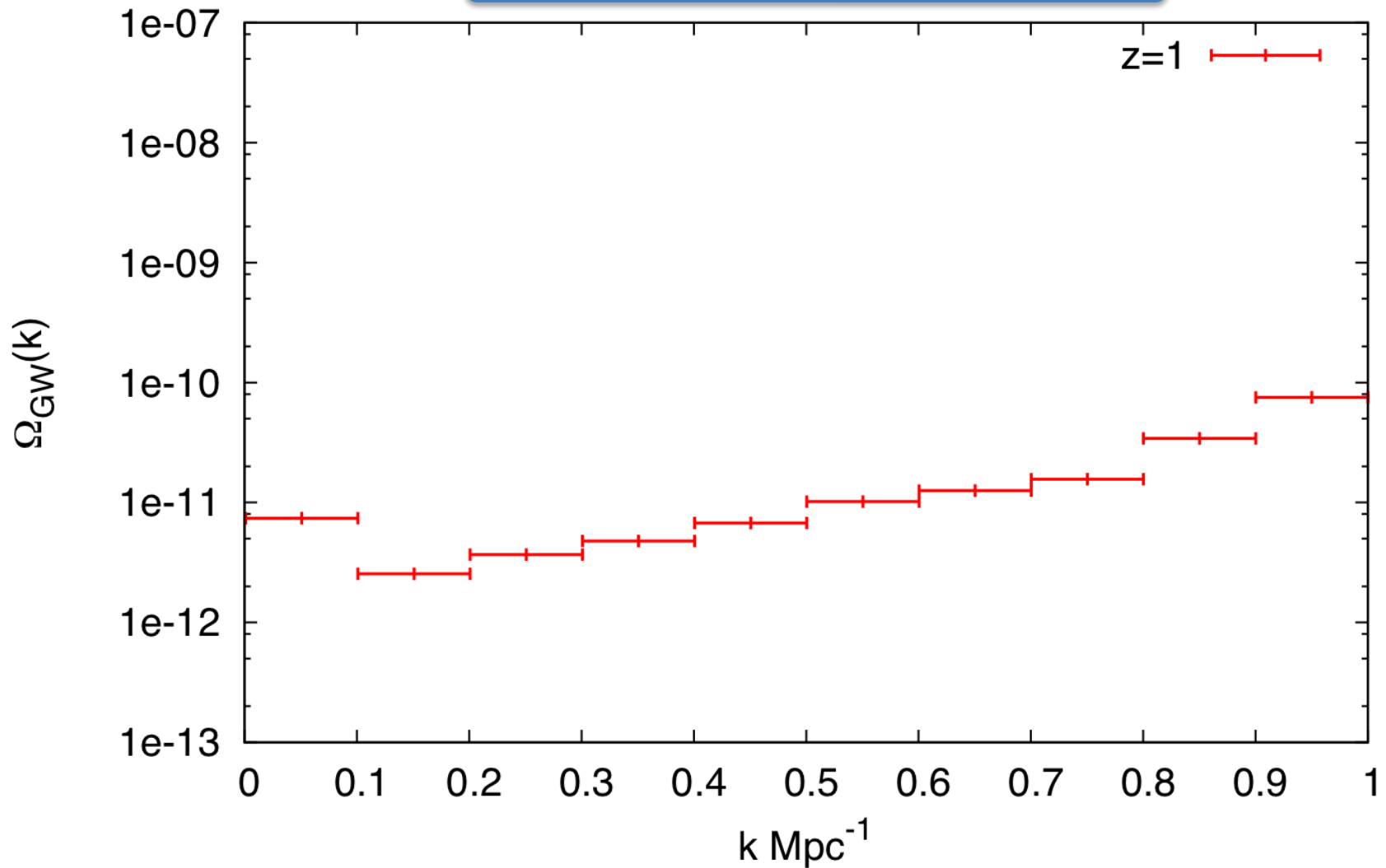


RMS deflection power is kept constant : $\langle \alpha^2 \rangle = \frac{1}{4\pi} \sum_l l(l+1)(2l+1) C_l^{h\otimes}$

$R = 1.9 \times 10^{-4}$
Scalar Perturbations

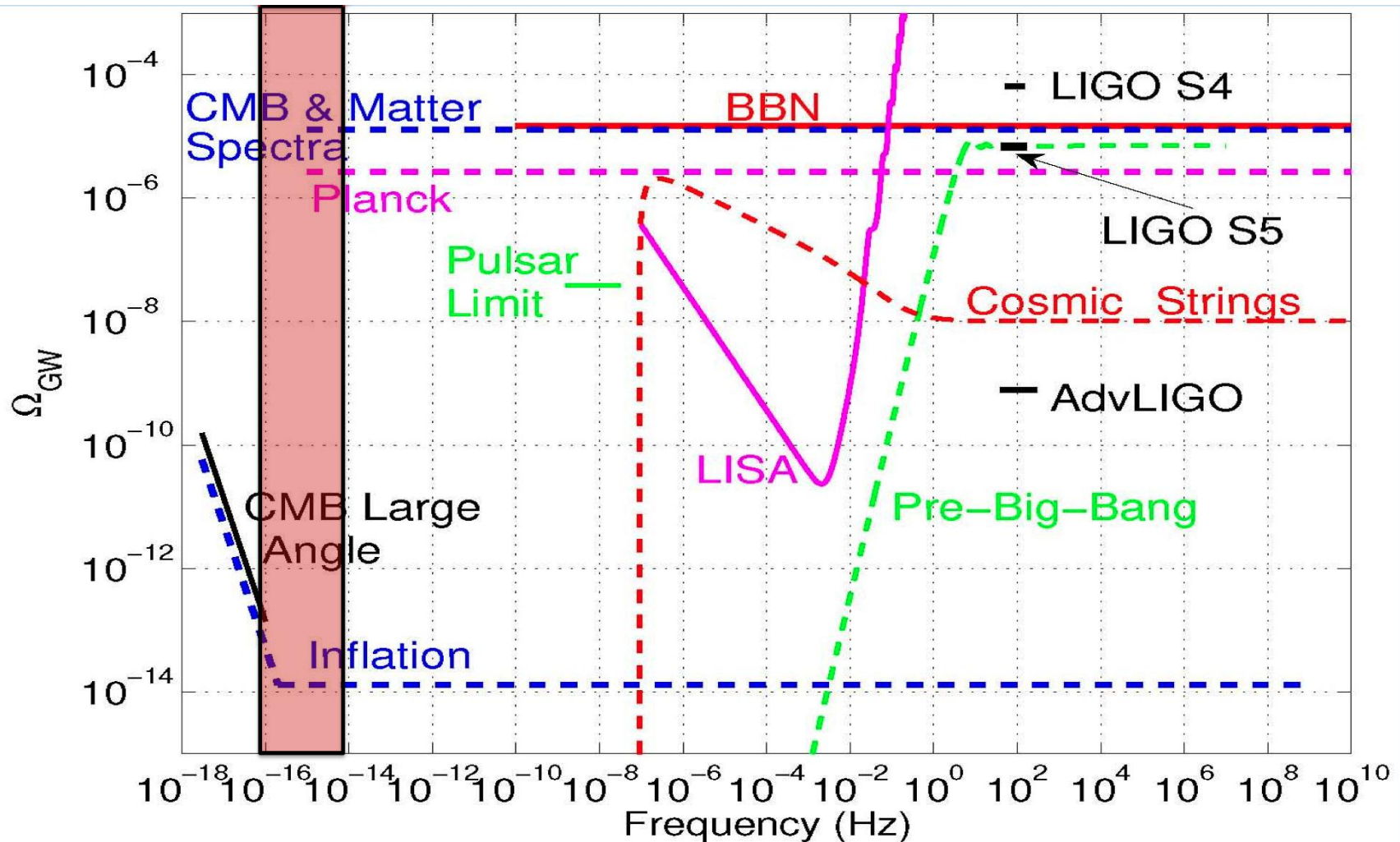
Constraints on the GW power spectra.

Preliminary Results



A New Window into Gravitational Waves !!

$k = 0.001 - 1 \text{ Mpc}^{-1}$ \longrightarrow $10^{-16} - 10^{-14} \text{ Hz}$



Summary :

- Tensors are more efficient at transferring power from EE \rightarrow BB spectra.
- From current and future experiments which will measure the CMB polarization spectra, it will be possible to put constraints on the Gravitational Waves spectra.
- This probe provides a new window into Gravitational Waves which has not been previously explored.

Thank you for your attention 😊