

Molecular hydrogen in high-redshift Damped Lyman- α systems: The VLT/UVES database^{*}

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ABSTRACT

Aims. We present the current status of ongoing searches for molecular hydrogen in high-redshift ($1.8 < z_{\text{abs}} \leq 4.2$) Damped Lyman- α systems (DLAs) capitalising on observations performed with the ESO Very Large Telescope (VLT) Ultraviolet and Visual Echelle Spectrograph (UVES).

Methods. We identify 77 DLAs/strong sub-DLAs, with $\log N(\text{H I}) \geq 20$ and $z_{\text{abs}} > 1.8$, which have data that include redshifted H_2 Lyman and/or Werner-band absorption lines. This sample of H I , H_2 and metal line measurements, performed in an homogeneous manner, is more than twice as large as our previous sample (Ledoux et al. 2003) considering every system in which searches for H_2 could be completed so far, including all non-detections.

Results. H_2 is detected in thirteen of the systems, which have molecular fractions of values between $f \approx 5 \times 10^{-7}$ and $f \approx 0.1$, where $f = 2N(\text{H}_2)/(2N(\text{H}_2) + N(\text{H I}))$. Upper limits are measured for the remaining 64 systems with detection limits of typically $\log N(\text{H}_2) \sim 14.3$, corresponding to $\log f < -5$. We find that about 35% of the DLAs with metallicities relative to solar $[\text{X}/\text{H}] \geq -1.3$ (i.e., $1/20^{\text{th}}$ solar), with $\text{X} = \text{Zn}$, S or Si , have molecular fractions $\log f > -4.5$, while H_2 is detected – regardless of the molecular fraction – in $\sim 50\%$ of them. In contrast, only about 4% of the $[\text{X}/\text{H}] < -1.3$ DLAs have $\log f > -4.5$. We show that the presence of H_2 does not strongly depend on the total neutral hydrogen column density, although the probability of finding $\log f > -4.5$ is higher for $\log N(\text{H I}) \geq 20.8$ than below this limit (19% and 7% respectively). The overall H_2 detection rate in $\log N(\text{H I}) \geq 20$ DLAs is found to be about 16% (10% considering only $\log f > -4.5$ detections) after correction for a slight bias towards large $N(\text{H I})$. There is a strong preference for H_2 -bearing DLAs to have significant depletion factors, $[\text{X}/\text{Fe}] > 0.4$. In addition, all H_2 -bearing DLAs have column densities of iron into dust grains larger than $\log N(\text{Fe})_{\text{dust}} \sim 14.7$, and about 40% of the DLAs above this limit have detected H_2 lines with $\log f > -4.5$. This demonstrates the importance of dust in governing the detectability of H_2 in DLAs. Our extended sample supports neither the redshift evolution of the detection fraction of H_2 -bearing DLAs nor that of the molecular fraction in systems with H_2 detections over the redshift range $1.8 < z_{\text{abs}} \leq 3$.

Key words. cosmology: observations – quasars: absorption lines – galaxies: ISM – ISM: molecules

1. Introduction

Damped Lyman- α systems (DLAs) were discovered in the seventies (e.g., Lowrance et al. 1972; Beaver et al. 1972; Carswell et al. 1975; Wright et al. 1979) and identified afterwards as redshifted damping absorptions from large column densities of neutral atomic hydrogen (Smith et al. 1979).

Numerous DLAs, with $N(\text{H I}) \geq 2 \times 10^{20}$ atoms cm^{-2} , have been discovered through large dedicated surveys (e.g., Wolfe et al. 1986) and more recently thanks to the huge number of quasar spectra available from the Sloan Digital Sky Survey (Prochaska et al. 2005). Because DLAs contain most of the neutral hydrogen available for star formation in the Universe (Wolfe et al. 1986; Lanzetta et al. 1991) and are associated with numerous metal absorption lines, they probably arise in the interstellar medium of protogalaxies, progenitors

of present-day galaxies (see, e.g., Wolfe & Prochaska 2000; Haehnelt et al. 2000; Wolfe et al. 2005). Our understanding of DLAs is mainly based on the study of low-ionisation metal absorptions (e.g., Prochaska & Wolfe 2002) but also high-ionisation species (Lu et al. 1996; Wolfe & Prochaska 2000; Fox et al. 2007a,b) and, in a few cases, molecular absorptions (e.g., Ledoux et al. 2003). The latter are not conspicuous however in contrast to what is seen in the Galaxy and, for a long time, only the DLA towards Q 0528–2505 was known to contain H_2 molecules (Levshakov & Varshalovich 1985). H_2 -bearing DLAs are nevertheless crucial to understand the nature of DLAs because molecular hydrogen is an important species to derive the physical conditions in the gas (see, e.g., Tumlinson et al. 2002; Reimers et al. 2003; Hirashita & Ferrara 2005; Srianand et al. 2005; Cui et al. 2005; Noterdaeme et al. 2007b).

The first systematic search for molecular hydrogen in high-redshift ($z_{\text{abs}} > 1.8$) DLAs was carried out using the Ultraviolet and Visual Echelle Spectrograph (UVES) at the Very Large Telescope (VLT) (Ledoux et al. 2003). It consisted of a sample of 33 DLAs with H_2 detected in eight of them. Molecular

^{*} Based on data gathered at the European Southern Observatory (ESO) using the Ultraviolet and Visual Echelle Spectrograph (UVES) installed at the Very Large Telescope (VLT), Unit-2 (Kueyen), on Cerro Paranal, Chile.

fractions were found to lie in the range $-3.5 < \log f < -1$ with $f = 2N(\text{H}_2)/(2N(\text{H}_2) + N(\text{H I}))$. Upper limits of typically $N(\text{H}_2) \sim 2 \times 10^{14} \text{ cm}^{-2}$ (corresponding to $\log f < -5$) were measured in the other systems. More recently, we noted a correlation between the presence of molecular hydrogen and the metallicity of high-redshift DLAs (Petitjean et al. 2006). High molecular fractions ($\log f > -4$) were found in about 40% of the high-metallicity DLAs ($[\text{X}/\text{H}] \geq 1/20^{\text{th}}$ solar) whilst only $\sim 5\%$ of the $[\text{X}/\text{H}] < -1.3$ DLAs have $\log f > -4$. Other papers by our group focused on specific detections. We presented the analysis of three systems with low molecular fractions, i.e., $\log f < -4$, one of them having a low metallicity (Noterdaeme et al. 2007a), and the H_2 -bearing DLA with, to date, the highest redshift, at $z_{\text{abs}} = 4.224$ towards Q 1441+2737 (Ledoux et al. 2006b). We note that a possible detection of H_2 in a DLA towards a Gamma-ray Burst (GRB) afterglow has been reported recently (Fynbo et al. 2006). However, the origin of DLAs at the GRB host-galaxy redshift is very likely to be different from those observed in QSO spectra (e.g., Jakobsson et al. 2006; Prochaska et al. 2007a).

We present here the whole sample of UVES high-redshift QSO-DLAs for which the wavelength range where H_2 lines are redshifted is covered by the available spectra. This sample is more than twice as large as in our previous study (Ledoux et al. 2003). We present the observations and the UVES DLA sample in Sect. 2 and provide comments on individual absorbers in Sect. 3. We discuss the overall population in Sect. 4 and results in Sects. 5 to 8. We conclude in Sect. 9.

2. Observations and sample

All the quasars in our sample were observed with the Ultraviolet and Visual Echelle Spectrograph (UVES, Dekker et al. 2000) mounted on the ESO VLT-UT2 (Kueyen) 8.2 m telescope on Cerro Paranal, Chile. We have used our UVES database to build up a sample of 77 DLAs/strong sub-DLAs along 65 lines of sight, hereafter called sample *S*. The systems were selected as having $N(\text{H I}) \geq 10^{20} \text{ cm}^{-2}$, and redshifts $z_{\text{abs}} > 1.8$ so that at least part of the wavelength range into which H_2 absorption lines are expected to be redshifted, is covered. Systems for which the flux at the wavelengths of all redshifted H_2 lines is zero as a consequence of the presence of an intervening Lyman-limit system located at a higher redshift were rejected. Systems for which the Ly- α forest is so crowded that no meaningful upper limit on $N(\text{H}_2)$ could be derived were also excluded from the final sample.

Sample *S* comprises a total of 68 DLAs ($\log N(\text{H I}) \geq 20.3$) and nine strong sub-DLAs ($20 \leq \log N(\text{H I}) < 20.3$) according to the common definition of DLAs ($\log N(\text{H I}) \geq 20.3$; Wolfe et al. 1986). Our lower limit on $N(\text{H I})$ for sub-DLAs is only slightly lower than the classical definition of DLAs ensuring that these absorbers are mostly neutral and share the same physical nature as classical DLAs (Viegas 1995).

Most of the systems in sample *S* (53 out of 77) come from the sample of Ledoux et al. (2006a). This sample is mainly drawn from the follow-up of the Large Bright QSO Survey (LBQS, Wolfe et al. 1995) and has been observed between 2000 and 2004 in the course of our systematic search for molecular hydrogen at $z_{\text{abs}} > 1.8$ (see also Petitjean et al. 2000; Ledoux et al. 2003). This comprises 46 bona-fide DLAs and seven strong sub-DLAs.

In addition to this sample, we are considering UVES data for 13 absorbers (11 DLAs and two strong sub-DLAs) from the Hamburg-ESO DLA survey (Smette et al. 2005), seven DLAs

from the CORALS survey (Akerman et al. 2005), and four DLAs mainly from our own observing runs. Among the latter systems, the $z_{\text{abs}} = 3.692$ and 3.774 DLAs towards Q 0131+0345 and the $z_{\text{abs}} = 2.659$ DLA towards Q 0642–5038 were observed in visitor mode on September 17–20 2004, under Prog. ID 073.A-0071 (PI: Ledoux). On the other hand, the $z_{\text{abs}} = 4.203$ DLA towards Q 0951–0450 was observed in service mode on January 26–27 and February 19–21, 2004, under Prog. ID 072.A-0558 (PI: Vladilo).

We have reduced all the data including those retrieved from the ESO archive in an homogeneous manner using the UVES pipeline (Ballester et al. 2000), which is available as a dedicated package of the ESO MIDAS data reduction system. The main characteristics of the pipeline are to perform a robust inter-order background subtraction for master flat-fields and science frames and an optimal extraction of the object signal subtracting the sky spectrum and rejecting cosmic rays simultaneously. The wavelength scale of the reduced spectra was converted to vacuum-heliocentric values. Each spectrum, corresponding to different instrument settings, was rebinned to a constant wavelength step. No further rebinning was performed during subsequent data analysis. Individual exposures were then weighted, scaled and combined altogether.

Total neutral hydrogen column densities have been measured from Voigt-profile fitting of the damped Lyman- α and/or Lyman- β lines (see Ledoux et al. 2006a; Smette et al., in prep). Most of the metallicity and depletion measurements are from Ledoux et al. (2006a) except when more recently re-measured (see details in footnote of Table 1). Apart from the DLAs towards Q 0131+0345 and Q 0951–0450 (for which abundances were taken from Prochaska et al. (2007b) because UVES observations did not allow a sufficient number of metal lines to be covered), we (re)determined all metal column densities by the homogeneous fitting of Voigt-profiles to non-saturated absorption lines. This includes the systems from the Hamburg-ESO and CORALS DLA surveys. Metallicities were determined using the zinc abundance (when Zn II is detected), or from the abundances of sulphur or silicon. No ionisation correction has been applied. Zinc and sulphur are known to be little depleted into dust grains. Silicon, in turn, is probably mildly depleted and has been used only in cases where Zn and S are not detected. The results are summarised in Table 1. Abundances are given relative to solar, i.e., $[\text{X}/\text{H}] \equiv \log N(\text{X})/N(\text{H}) - \log (\text{X}/\text{H})_{\odot}$. Solar abundances as listed in Morton (2003), based on meteoritic data from Grevesse & Sauval (2002), were adopted.

3. Comments on individual systems

Q 0000–2619, $z_{\text{abs}} = 3.390$: Levshakov et al. (2000, 2001) reported a tentative detection of H_2 in this system. However, they detected only two weak absorption features in the Lyman- α forest, identified as H_2 W2-0 Q1 and L4-0 R1. The probability that these features are actually due to intervening Lyman- α absorbers is high and we consider the derived $J = 1$ column density as an upper limit only. In any case, this system has a very low molecular fraction of $\log f < -6.87$.

Q 0013–0029, $z_{\text{abs}} = 1.973$: The detection of H_2 in this system has been reported for the first time by Ge & Bechtold (1997). Petitjean et al. (2002) showed that this system is actually the blend of a DLA ($\log N(\text{H I}) = 20.83$) and a sub-DLA ($\log N(\text{H I}) \leq 19.43$) separated by $\sim 500 \text{ km s}^{-1}$. H_2 is present in both systems, with four components detected up to $J = 5$.

Q 0027–1836, $z_{\text{abs}} = 2.402$: This system has the lowest metallicity ($[\text{X}/\text{H}] = -1.63$) of systems in which H_2 molecules have

been detected. H_2 is detected in a single component up to rotational level $J = 5$ and possibly 6. Thanks to the high data quality, we detected an increase in the Doppler parameter b , from low to high rotational levels, for the first time at high redshift (Noterdaeme et al. 2007a).

Q 0347–3819, $z_{\text{abs}} = 3.025$: The detection of H_2 in this system was first reported by Levshakov et al. (2002) using UVES commissioning data. Subsequently Ledoux et al. (2003) analysed higher quality spectra of the system. Weak molecular absorption is observed in a single component up to $J = 4$.

Q 0405–4418, $z_{\text{abs}} = 2.595$: Molecular hydrogen is detected up to $J = 3$ in a single component (see Ledoux et al. 2003; Srianand et al. 2005). There is some indication that the relative strengths of the $J = 2$ and 3 lines require $b \gtrsim 1.5 \text{ km s}^{-1}$, while the $J = 0$ and 1 lines are consistent with $b = 1.3 \text{ km s}^{-1}$.

Q 0528–2505, $z_{\text{abs}} = 2.811$: This is the first DLA system in which H_2 was detected (Levshakov & Varshalovich 1985). The detection of H_2 was readdressed in more detail by Foltz et al. (1988), Srianand & Petitjean (1998) and Srianand et al. (2005). The single H_2 component seen in a low resolution CASPEC spectrum (Srianand & Petitjean 1998) is resolved into two components in a UVES spectrum (Srianand et al. 2005). Rotational levels up to $J = 5$ are detected.

Q 0551–3638, $z_{\text{abs}} = 1.962$: This system has the highest metallicity in the UVES DLA sample with $[\text{Zn}/\text{H}] = -0.35$. H_2 is clearly detected in two well-resolved features separated by about 55 km s^{-1} (Ledoux et al. 2002). The first feature is narrow and weak, but detected up to $J = 3$. The second feature is broader and stronger and has been fitted using two components at precisely the same redshift as the detected C I lines.

Q 0642–5038, $z_{\text{abs}} = 2.659$: H_2 in this system is detected in a single strong component. This is a new detection and details will be presented in a forthcoming paper (Ledoux et al., in prep.).

Q 0841+1256, $z_{\text{abs}} = 2.375$: Petitjean et al. (2000) detected two weak absorption features corresponding to H_2 L4-0R0 and L2-0R0 at the same redshift. As for the DLA towards Q 0000–2619, we consider the tentative measurement of the $J = 0$ column density as an upper limit only.

Q 1232+0815, $z_{\text{abs}} = 2.338$: A single H_2 component is seen in rotational levels $J = 0$ to 5 (Ge et al. 2001; Srianand et al. 2005). This is the only DLA in which deuterated molecular hydrogen (HD) has been detected until now (Varshalovich et al. 2001).

Q 1441+2737, $z_{\text{abs}} = 4.224$: Because of the high redshift of this system (actually the highest amongst known H_2 -bearing DLAs), the Lyman- α forest is quite dense in the wavelength region in which H_2 lines are redshifted. However, H_2 could be detected and studied in a large number of transitions thanks to the high spectral resolution of the data. Four rotational levels are detected in no less than three velocity components (Ledoux et al. 2006b).

Q 1444+0126, $z_{\text{abs}} = 2.087$: Molecular hydrogen is found in two components of this sub-DLA in $J \leq 3$ rotational levels (Ledoux et al. 2003).

Q 1451+1223, $z_{\text{abs}} = 2.469$: Because of the Lyman-break from the DLA at $z_{\text{abs}} = 3.171$, the expected positions of only two lines (H_2 L0-0R0 and L0-0R1) fall in a region of non-zero flux. The high upper limit we derive is not strongly constraining and the molecular fraction is probably well below $\log f = -4.5$.

Q 2318–1107, $z_{\text{abs}} = 1.989$: Weak H_2 absorption features from $J = 0$ to 2 are unambiguously detected in a single component (Noterdaeme et al. 2007a).

Q 2343+1232, $z_{\text{abs}} = 2.431$: Thanks to the high quality of the spectrum, H_2 absorption lines (from $J = 0$ and 1) are unambiguously detected although they are extremely weak ($N(\text{H}_2) \sim 5 \times 10^{13} \text{ cm}^{-2}$). We measure the lowest molecular fraction yet

measured for a DLA, $\log f = -6.41$ (Petitjean et al. 2006; Noterdaeme et al. 2007a).

Q 2348–0108, $z_{\text{abs}} = 2.426$: No less than seven H_2 components are detected in this system (Petitjean et al. 2006; Noterdaeme et al. 2007b). The H_2 profile is complex with three strong and four weak components spread over about 250 km s^{-1} , making H_2 lines from different transitions overlap. However, the large number of observed transitions and the good data quality allow an accurate measurement of column density for different rotational levels.

4. Overall population

In this Section, we use the results summarised in Table 1 to derive the properties of the H_2 gas in the global DLA population. We list in this table, the name of the quasar in column 1, the emission and absorption redshifts in columns 2 and 3, the total H I column density in column 4, the metallicity relative to solar $[\text{X}/\text{H}]$ and the depletion factor $[\text{X}/\text{Fe}]$ in columns 5 and 6, with X given in column 7, the velocity spread of the low-ionisation metal line profiles (see Ledoux et al. 2006a) in column 8, the total molecular hydrogen column densities in the $J = 0$ and 1 rotational levels in columns 9 and 10, the mean molecular fraction in column 11, and the references for the measurements in column 12. Upper limits on the molecular fraction were calculated to be the sum of the upper limits on the first two rotational levels ($J = 0$ and 1) and are given at the 3σ significance level. The thirteen firm detections of H_2 constitute what we will hereafter call sub-sample \mathcal{S}_{H_2} . Ten of them have molecular fractions higher than all upper limits, i.e., $\log f > -4.5$, and will constitute sub-sample \mathcal{S}_{HF} (standing for high molecular fractions). We note that obviously $\mathcal{S}_{\text{HF}} \subset \mathcal{S}_{\text{H}_2} \subset \mathcal{S}$.

The upper panel of Fig. 1 gives the distribution of $N(\text{H I})$ from sample \mathcal{S} , sub-samples \mathcal{S}_{H_2} and \mathcal{S}_{HF} to be compared to the reference distribution derived from the SDSS-DR5 DLA sample scaled to the number of systems in the UVES sample (Prochaska et al. 2005). It is apparent that the distribution from the UVES sample is slightly biased compared to the scaled SDSS distribution¹ in favour of large $N(\text{H I})$ DLAs (i.e., with $\log N(\text{H I}) \geq 20.8$). We can however correct for this bias. We divide each DLA sample into two halves, above and below $\log N(\text{H I}) = 20.8$. We scale the number of systems in each bin of the UVES distribution by the ratio of the number of systems in the (scaled) SDSS and UVES samples, above (by a ratio of ~ 0.5) and below (by ~ 1.4) $\log N(\text{H I}) = 20.8$. The result of this scaling is shown in the bottom panel of Fig. 1. In the inset, the cumulative distributions are shown in order to compare the two populations. It is apparent that the scaling has corrected for the above-mentioned bias.

The overall H_2 detection rate in $\log N(\text{H I}) \geq 20$ systems is found to be about 16% (10% considering only $\log f > -4.5$ detections) after correction for the above-mentioned bias. There is no H_2 detection for $\log N(\text{H I}) < 20.2$. These numbers, however, may well be affected by small number statistics. H_2 is in fact detected in a sub-DLA component ($\log N(\text{H I}) \sim 19.4$) of the $z_{\text{abs}} = 1.973$ system towards Q 0013–0029 (see Petitjean et al. 2002). The corresponding molecular fraction is larger than 10^{-4} . In any case, additional observations are required to investigate the molecular content of sub-DLAs. Apart, possibly, from the first bin ($\log N(\text{H I}) < 20.2$), there is no strong dependence of

¹ Note that the SDSS $N(\text{H I})$ distribution itself may be biased, however it is currently the best available representation of the DLA population.

the presence of molecules on the neutral hydrogen column density. Large molecular fractions are seen over the range $20.2 \leq \log N(\text{H I}) < 21.8$.

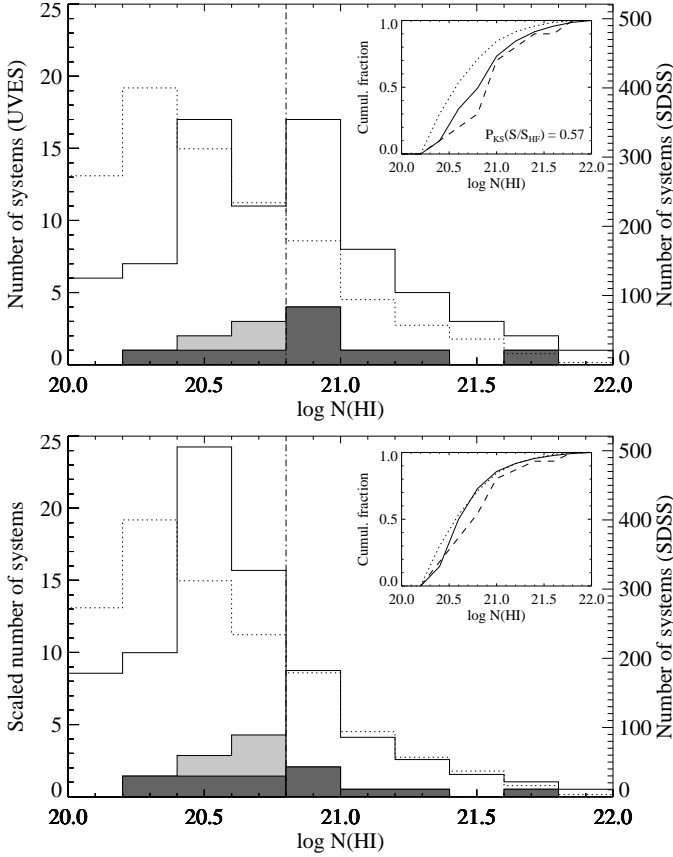


Fig. 1. *Top panel:* Neutral hydrogen column density distributions of DLAs in the overall UVES sample S (solid), the sub-sample of H_2 -bearing systems S_{H_2} (grey), and that of the systems with $\log f > -4.5$ (S_{HF} ; dark grey). The distribution from the SDSS-DR5 sample (dotted; Prochaska et al. 2005) is represented with a different scaling (right axis) so that the area of both histograms are the same. The inset shows the cumulative distributions from sample S (solid line), sub-sample S_{HF} (dashed), and the SDSS-DR5 sample (dotted). *Bottom panel:* The above $\log N(\text{H I})$ distributions are corrected for the bias in favour of large $N(\text{H I})$ column densities (see text) by scaling the numbers separately for $20 \leq \log N(\text{H I}) < 20.8$ and $\log N(\text{H I}) \geq 20.8$ to match the SDSS-DR5 distribution.

The double-sided Kolmogorov-Smirnov test yields a probability $P_{\text{KS}} = 0.57$ that sample S and sub-sample S_{HF} are drawn from the same parent population.

However, molecular hydrogen with $\log f > -4.5$ is detected in respectively 9% and 19% of the systems in the two sub-samples: $20.2 \leq \log N(\text{H I}) < 20.8$ (3/35) and $\log N(\text{H I}) \geq 20.8$ (7/36), which implies that there is probably a tendency for H_2 to be seen more often in large column density systems. We note that the fraction is 7% (3/41) for the range $20.0 \leq \log N(\text{H I}) < 20.8$. We also note that we are considering here the total neutral hydrogen column densities. The actual H I -clouds corresponding to the H_2 absorptions probably have smaller $N(\text{H I})$. This is in fact observed for the DLA at $z_{\text{abs}} = 1.973$ towards Q 0013–0029, for which partial deblending of the Ly- α absorption is possible.

The fact that the sample is biased in favour of large H I column density DLAs could be a problem for the analysis of the overall DLA population if any correlation exists between $\log N(\text{H I})$ and the metallicity ($[\text{X}/\text{H}]$) and/or the depletion of metals onto dust grains ($[\text{X}/\text{Fe}]$). This is not the case however: Figure 2 shows the neutral hydrogen column density as a function of metallicity (top panel) and depletion factor (bottom panel). It is apparent that there is no correlation between $[\text{X}/\text{H}]$ or $[\text{X}/\text{Fe}]$ and $\log N(\text{H I})$. The Kendall's rank correlation coefficient is as small as 0.02 (resp. -0.01) for $[\text{X}/\text{H}]$ (resp. $[\text{X}/\text{Fe}]$) vs. $\log N(\text{H I})$. The lack of systems with both a high metallicity and large $N(\text{H I})$ is a feature common to all DLA samples. It can be explained either by a lack of large H I column densities because of the transition of H I into H_2 (Schaye 2001) and/or by the fact that any quasar located behind such an absorber would remain undetected because of the induced extinction (e.g., Boissé et al. 1998; Vladilo & Péroux 2005). Such systems could also be associated with regions with low projected cross-sections. Therefore, although the lack of high-metallicity systems with large $N(\text{H I})$, together with the bias towards large $N(\text{H I})$ in our sample, could introduce a small bias towards low metallicities, the absence of the above-described correlations implies that any bias will have little influence on the properties based on the metallicities and depletion factors discussed in Sect. 6.

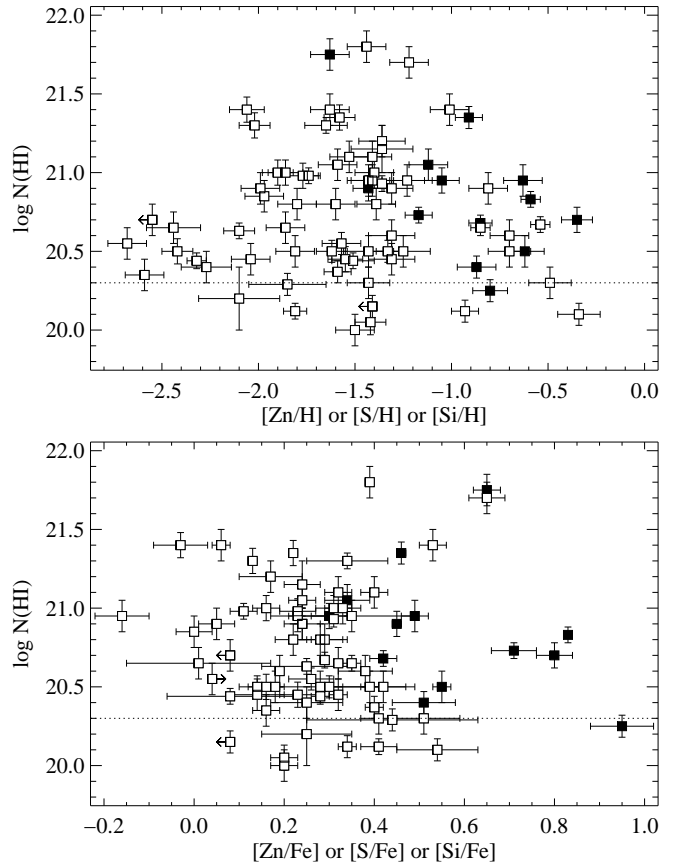


Fig. 2. Logarithm of the total neutral hydrogen column density in the UVES sample versus metallicity (top panel) or depletion factor (bottom panel). Filled squares indicate systems in which H_2 is detected. The lack of systems with both a high metallicity and a large H I column density is seen in all DLA samples. There is no correlation between $N(\text{H I})$ and $[\text{X}/\text{H}]$ or $[\text{X}/\text{Fe}]$.

In Fig. 3, we compare the metallicity distribution of DLAs in the UVES sample to that of DLAs in the Keck sample (HIRES+ESI, Prochaska et al. 2007b). Although a difference can be seen in the range $-1.6 < [X/H] < -1$, partly due to the bias towards large $N(\text{H I})$, no systematic bias is observed for the UVES sample, apart from any also affecting the Keck sample, as shown by the cumulative distributions (inset). We note that the observed distributions are similar in shape to the distribution derived using simulations by Hou et al. (2005, their fig. 3). The metallicity distribution for H_2 -bearing DLAs is also plotted. It is apparent from the figure that the distribution of H_2 -bearing systems is strongly skewed towards high metallicities (see Sect. 6).

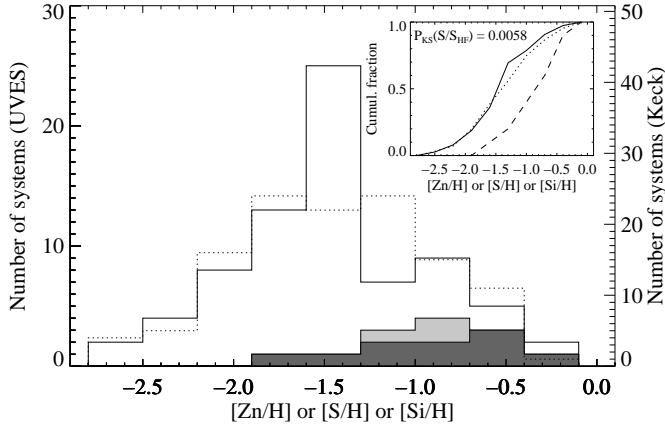


Fig. 3. The metallicity distribution of DLAs in the overall UVES sample (\mathcal{S} ; solid) is compared to that from the Keck sample (dotted; Prochaska et al. 2007b), adequately scaled (right axis) so that the area of both histograms are the same. The distributions are similar. Distributions from sub-samples \mathcal{S}_{H_2} (grey) and \mathcal{S}_{HF} (dark grey) are also shown. It is clear from this and the Kolmogorov-Smirnov test probability (inset) that the distributions from H_2 -bearing systems are skewed towards high metallicities.

5. Molecular fraction

As can be seen in Fig. 4, there is no sharp transition in the molecular fraction of DLAs at any total hydrogen ($\text{H I} + \text{H}_2$) column density $N(\text{H})$, in contrast to what is observed at $\log N(\text{H}) = 20.7$ in the Galactic disk (Savage et al. 1977, bottom panel of Fig. 4) or at $\log N(\text{H}) = 20.4$ along high-latitude Galactic lines of sight (Gillmon et al. 2006). We note also that no sharp transition from H I to H_2 is observed in the Magellanic clouds (upper and middle panels of Fig. 4), although SMC lines of sight probe larger gas column densities. Such transitions are expected to occur at $\log N(\text{H}) > 21.3$ in the LMC and $\log N(\text{H}) > 22$ in the SMC (Tumlinson et al. 2002). The molecular fractions of DLAs are similar to values measured for the Magellanic clouds, but are lower than measurements for the Galactic disk. The absence of a sharp $\text{H I}/\text{H}_2$ transition does not contradict the predictions by Schaye (2001) because the covering factor of large $N(\text{H}_2)$ systems is expected to be small (Zwaan & Prochaska 2006). Moreover, the induced extinction could be large (see below) and such systems may have so far been missed. We should compile a representative sample of strong DLAs ($\log N(\text{H I}) > 21.5$) in order to address this point.

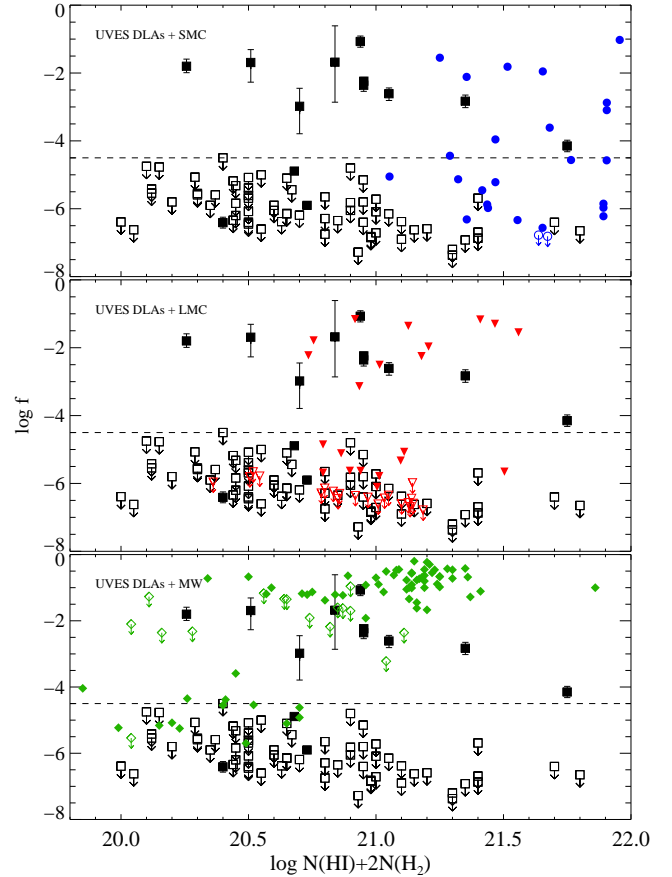


Fig. 4. Molecular fraction versus total hydrogen ($\text{H I} + \text{H}_2$) column density for sample \mathcal{S} (the filled squares correspond to H_2 detections, i.e., sub-sample \mathcal{S}_{H_2}), lines of sight through the SMC (blue dots, top panel; Tumlinson et al. 2002), the LMC (red triangles, middle panel; Tumlinson et al. 2002) and the Galactic disk (green diamonds, bottom panel; Savage et al. 1977). Upper limits are marked by unfilled symbols with downward arrows.

6. The importance of dust

Figure 5 is an extension of fig. 3 from Petitjean et al. (2006) to the whole UVES sample, with about 2.5 times more systems with $[X/H] < -1.3$. It can be seen from this figure that about 35% of the $[X/H] \geq -1.3$ ($1/20^{\text{th}}$ solar) systems have molecular fractions $\log f > -4.5$, while H_2 is detected – in $\sim 50\%$ of them. In contrast, only about 4% of the $[X/H] < -1.3$ DLAs have $\log f > -4.5$. We remind the reader that we use the limit $\log f > -4.5$ to define systems with high molecular fraction because all of our upper limits are below this value. The lowest metallicity at which H_2 has been detected to date in a DLA is $[\text{Zn}/\text{H}] = -1.63$ (Noterdaeme et al. 2007a). The corresponding system, at $z_{\text{abs}} = 2.402$ towards Q 0027–1836, has a molecular fraction of $\log f = -4.15$.

This trend in metallicity supports the idea that the presence of dust is an important ingredient in the formation of H_2 . The correlation between the depletion of metals into dust grains and the metallicity (Ledoux et al. 2003) has now been confirmed using a sample more than twice the size, and at higher confidence level (5.3σ compared to 4σ) (see Fig. 6). High-metallicity DLAs are usually more dusty and, as a consequence, the for-

mation rate of H_2 onto dust grains is enhanced, while the photo-destruction rate is lowered by dust- and self-shielding.

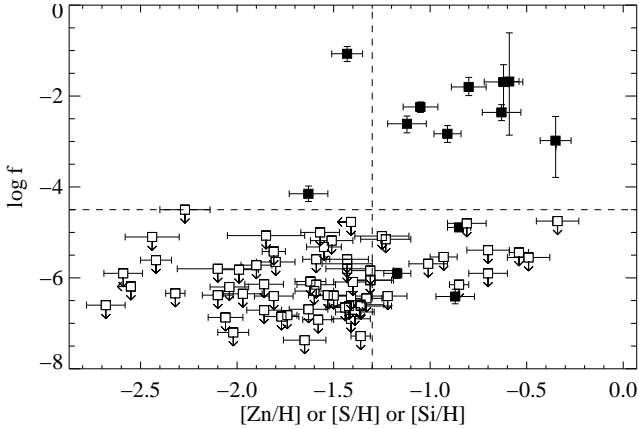


Fig. 5. Logarithm of the molecular fraction, $f = 2N(H_2)/(2N(H_2) + N(HI))$, versus metallicity, $[X/H] = \log N(X)/N(H) - \log(X/H)_\odot$ with $X=Zn, S$ or Si . Filled squares indicate systems in which H_2 is detected. The horizontal (resp. vertical) dashed line indicates the separation between what we somewhat arbitrarily call high and low molecular fractions, $\log f = -4.5$ (resp. high and low metallicities, $[X/H] = -1.3$).

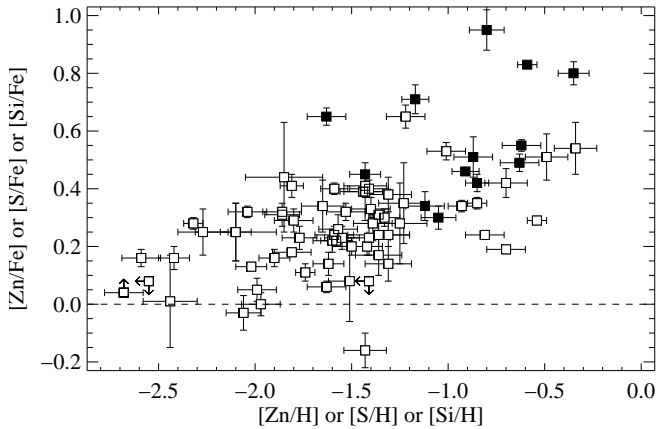


Fig. 6. A correlation (5.3σ confidence level) exists in the overall UVES sample between dust-depletion factor, $[X/Fe]$, and metallicity, $[X/H]$, with X either Zn, S or Si .

Figure 7 shows the distribution of depletion factors (defined as $[X/Fe]$, with $X=Zn, S$ or Si) for the different samples of systems defined in this paper. It is clear from this histogram that H_2 is found in DLAs with the highest depletion factors. The probability that sample S and sub-sample S_{H_2} are drawn from the same parent population is very small, i.e., $P_{KS} \approx 10^{-3}$.

To test the influence of the presence of dust on the molecular hydrogen column density, it is worth comparing $N(H_2)$ to the column density of iron into dust, $N(Fe)_{dust} = (1 - 10^{-[X/Fe]})N(X)(Fe/X)_{dla}$ (see Fig. 8). Vladilo et al. (2006) showed that absorbers associated with large column densities of iron in dust, generally induce more extinction (i.e. have larger

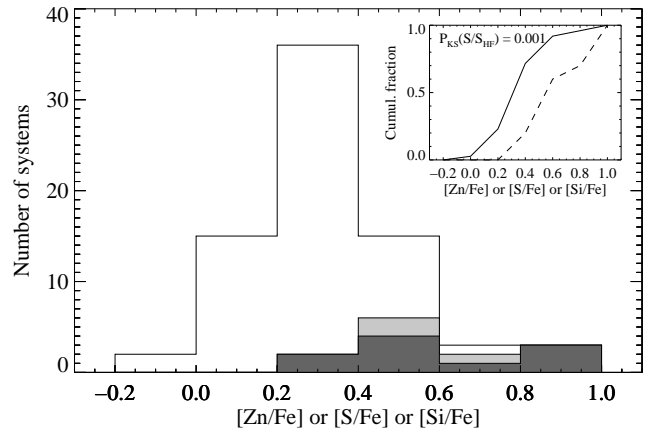


Fig. 7. Distribution of depletion factors from the overall UVES sample (S ; solid), sub-samples S_{H_2} (grey) and S_{HF} (dark grey). It is clear from these histograms that the distribution of depletion factors for H_2 -bearing DLAs is different from that of the overall sample. This shows clearly that H_2 -bearing DLAs are more dusty than the rest of the DLAs. The probability that the two samples are drawn from the same parent population is indeed very small: $P_{KS}(S/S_{HF}) \approx 10^{-3}$.

values of A_V). This correlation is consistent with that observed along Galactic interstellar lines of sight (e.g., Snow et al. 2002). We note that the above expression of $N(Fe)_{dust}$ assumes that X is a non-refractory element and that the intrinsic $[X/Fe]$ ratio is solar. Vladilo (2002) introduced corrections for zinc depletion and possible non-solar intrinsic abundances (i.e., $(Fe/X)_{dla} \neq (Fe/X)_\odot$; see also Vladilo et al. 2006) that we do not apply here because the corrections are small, and not necessary for our purposes. This effect is most important for Silicon that can be depleted into dust grains, and the depletion factor is underestimated when using $[Si/Fe]$. The effect is however probably less than 0.3 dex (see Petitjean et al. 2002) and should have no consequence on the characteristics of the whole population. When $[X/Fe]$ is negative, which happens for three DLAs, $N(Fe)_{dust}$ cannot be computed directly. We thus estimated a 3σ upper limit on $N(Fe)_{dust}$ by considering the upper bound provided by the 3σ error on $[X/Fe]$.

It can be seen on Fig. 8 that all H_2 -bearing DLAs have $N(Fe)_{dust} > 5 \times 10^{14} \text{ cm}^{-2}$, supporting the idea that dust is an important ingredient in the formation of H_2 . We note however that this column density corresponds to a small extinction ($\log A_V \sim -1.5$; Vladilo et al. 2006) possibly explaining the absence of detectable H_2 in three DLAs at $N(Fe)_{dust} > 3 \times 10^{15} \text{ cm}^{-2}$.

An additional reason why H_2 is not detected in these systems could be related to the particle density in the neutral gas being too small. The equilibrium between formation and destruction of H_2 molecules can be written as $R_{nn}(H) = R_{diss}n(H_2)$, where R_{diss} is the photo-dissociation rate, R is the formation rate, n the particle density, $n(H) \simeq n$ the proton density and $n(H_2)$ the H_2 density. By multiplying this expression by the longitudinal size of the cloud, one can see that the H_2 column density depends linearly on the particle density. Large dust column densities may occasionally result from large and diffuse clouds where the particle density is insufficiently large for H_2 to form efficiently enough to be detectable. The column densities of iron in dust are in fact large ($N(Fe)_{dust} > 3 \times 10^{15} \text{ cm}^{-2}$) in three DLAs without detectable H_2 mainly because of the large neu-

tral hydrogen column densities (i.e. $\log N(\text{H I}) = 21.70$ towards Q 0458–0203; $\log N(\text{H I}) = 21.80$ towards Q 1157+0128; and $\log N(\text{H I}) = 21.40$ towards Q 1209+0919).

The fact that H_2 is detected in some DLAs at low extinction, could imply that the weak (due to high depletion) iron component, associated with the H_2 -bearing component, is hidden in the overall metal-line profile (see, e.g., Ledoux et al. 2002). This may bias the measurements of metallicities and depletion factors, and hence the measurement of column density of iron into dust grains. In the future, we propose that this issue is more deeply investigated.

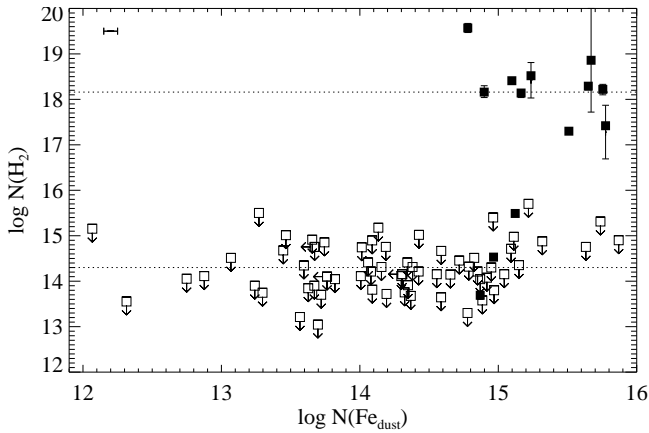


Fig. 8. H_2 column density, $N(\text{H}_2)$, versus column density of iron into dust grains, $N(\text{Fe})_{\text{dust}}$. The horizontal dotted line at $\log N(\text{H}_2) \approx 18.2$ (resp. 14.3) shows the median $\log N(\text{H}_2)$ (resp. the median upper limit on $\log N(\text{H}_2)$) for the H_2 -bearing DLAs from sub-sample \mathcal{S}_{H_2} (resp. the H_2 -undetected systems from sub-sample $\mathcal{S} - \mathcal{S}_{\text{H}_2}$). There is a clear gap of about four orders of magnitude between the two values. The horizontal bar in the upper-left corner of the plot shows the typical uncertainty on the measured $\log N(\text{Fe})_{\text{dust}}$.

7. Gas kinematics

Figure 9 shows the distribution of the velocity widths of the low-ionisation metal line profiles. These measurements were completed using the method described in Ledoux et al. (2006a), and many were previously published in that paper. It is clear from Fig. 9 that the probability of finding H_2 is higher when the velocity width of the low-ionisation metal line profiles is larger. Using the double-sided Kolmogorov-Smirnov test, we calculate a probability $P_{\text{KS}} = 0.075$ that samples \mathcal{S} and \mathcal{S}_{HF} are derived from the same parent distribution. A natural explanation is provided by Ledoux et al. (2006a) who suggested that higher metallicity DLAs arise from more massive objects. The amount of molecular gas and the star-formation rate in these systems could therefore be enhanced (Hirashita & Ferrara 2005) naturally explaining the above correlation. Another explanation could be that outflows can pull out cold and dusty gas, as well as providing a large number of velocity components that would increase the probability of finding H_2 (Murphy et al. 2007). However, H_2 is usually detected in few components (typically one or two) in the main clumps of the systems, and not in the satellite components at high velocities that make up most of the width of the profile. The only exceptions to this are the DLA

at $z_{\text{abs}} \approx 2.43$ towards Q 2348–0108 where no less than seven H_2 components are spread over 250 km s^{-1} (Noterdaeme et al. 2007b), and the $z_{\text{abs}} = 1.973$ system towards Q 0013–0029 where two pairs of H_2 components are separated by more than 500 km s^{-1} (Petitjean et al. 2002). The latter could be considered as the blend of two DLA systems (see Sect. 3). In addition, Fox et al. (2007a) interpreted DLAs where the C IV total line width exceeds the escape velocity to be associated with outflowing winds. Using this criterion, there is no indication that H_2 is found preferably in DLAs with outflows. Indeed, considering the 54 DLA systems common to the present sample and that of Fox et al. (2007a), we can estimate that H_2 is detected in $\sim 16\%$ (3/19) of the DLAs with outflows while this percentage is $\sim 22\%$ (12/54) for the entire DLA population.

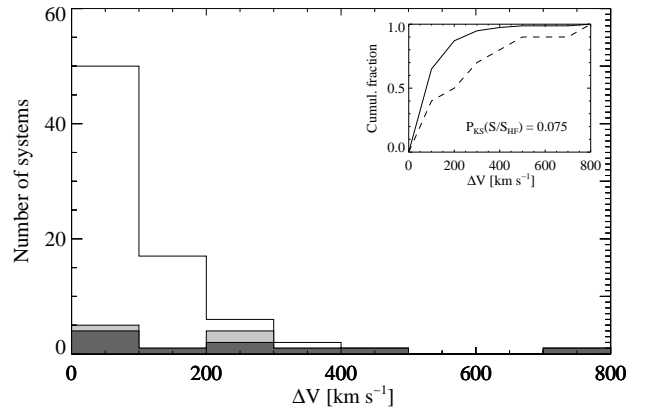


Fig. 9. Distribution of the velocity spread, ΔV , of low-ionisation metal line profiles (see Ledoux et al. 2006a, for details on the technique) from the overall UVES sample (\mathcal{S} , 77 systems, solid histogram), systems with detected molecular hydrogen (\mathcal{S}_{H_2} , 13 systems, grey), and systems with $\log f > -4.5$ (\mathcal{S}_{HF} , ten systems, dark grey). It is apparent from this and the Kolmogorov-Smirnov test probability (inset) that the likelihood of finding H_2 is higher when ΔV is larger.

8. Evolution with redshift

In Fig. 10, we analyse the redshift distribution of DLAs within the UVES sample. The H_2 -bearing DLA sub-sample is statistically indistinguishable from the overall UVES sample. A double-sided Kolmogorov-Smirnov test shows that sample \mathcal{S} and sub-sample \mathcal{S}_{H_2} have a probability $P_{\text{KS}} = 0.94$ to be drawn from the same parent population. There is no indication of any evolution of the fraction of H_2 -bearing DLAs with redshift. Molecular hydrogen is more difficult to detect at the lowest redshifts observable using UVES ($z_{\text{abs}} \lesssim 2$) because fewer and weaker transitions are then covered by the UVES spectra and lines are located in the bluest part of the spectra where the signal-to-noise ratios are lower. While this could imply a bias in the detection rate of H_2 in systems with low molecular fractions, high molecular fractions (i.e., $\log f > -4.5$) are detected anyway because the corresponding column densities are always well above the detection limit. When considering only high molecular fractions, the same Kolmogorov-Smirnov test yields $P_{\text{KS}}(\mathcal{S}/\mathcal{S}_{\text{HF}}) = 0.98$. Contrary to previous tentative evidence (Curran et al. 2004), there is also no evolution of the molecular fraction in systems with detected H_2 (see Fig. 10).

However, because of small number statistics for systems at $z > 3$, it is still unclear whether there is no really evolution in the sense of fewer H_2 detections at higher redshift. Therefore, we restrict our claim of no evolution with redshift of the detection fraction and the molecular fraction to the range $1.8 < z_{\text{abs}} \leq 3$. On the other hand, large molecular fractions are observed in the local Universe and significant evolutions of both the fraction of H_2 -detected systems and the molecular fractions are expected at lower redshifts. It is then of prime importance to cover the redshift interval $0 \leq z_{\text{abs}} \leq 1.8$. This will soon be possible using the Cosmic Origins Spectrograph to be installed onboard the Hubble Space Telescope.

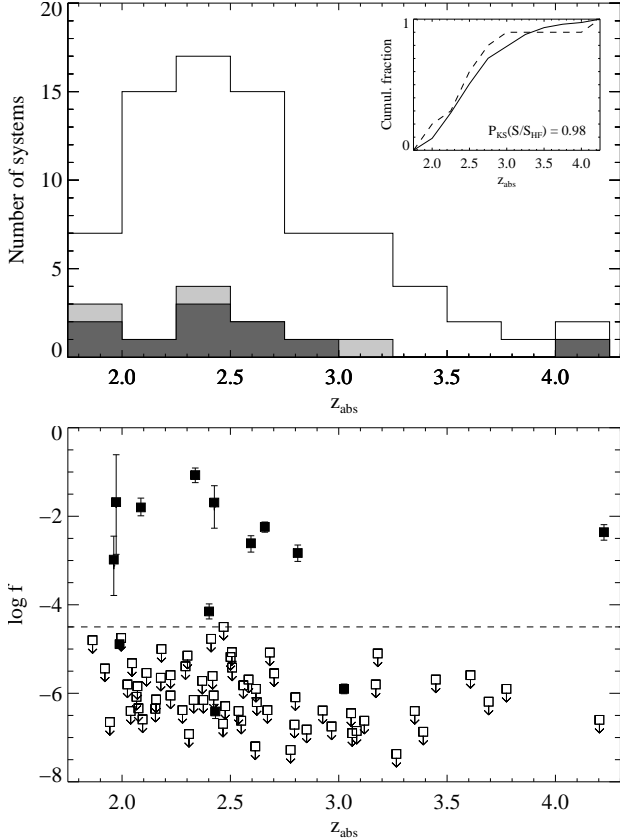


Fig. 10. *Top panel:* Redshift distributions of DLAs in the overall UVES sample (S , 77 systems, solid histogram), systems with detected molecular hydrogen (S_{H_2} , 13 systems, grey), and systems with $\log f > -4.5$ (S_{HF} , ten systems, dark grey). The inset shows the cumulative distributions from sample S (solid) and sub-sample S_{HF} (dashed). The two distributions are statistically indistinguishable ($P_{\text{KS}} = 0.98$). *Bottom panel:* Logarithm of the molecular fraction, $\log f$, versus absorber redshift. There is a slight trend for more constraining upper limits at higher redshift. This is not physical however but due to the higher signal-to-noise ratios achieved towards the red, together with the larger number of H_2 transitions covered by the spectra.

9. Conclusion

We present results of the largest survey of molecular hydrogen in high-redshift ($1.8 < z_{\text{abs}} \leq 4.2$) DLAs compiled to date using high signal-to-noise ratio, high spectral-resolution VLT-

UVES data. We analyse data for 77 DLAs/strong sub-DLAs with $N(\text{H I}) \geq 10^{20} \text{ cm}^{-2}$, a dataset more than twice as large as that studied previously by Ledoux et al. (2003). From the thirteen high-redshift H_2 -bearing DLAs known to date, nine have been discovered by our group. Due to the superb quality of the Ultraviolet and Visual Echelle Spectrograph, we are able to detect unambiguously the H_2 absorption features, measure accurate column densities, and in the cases of non-detections derive stringent upper limits.

A double-sided Kolmogorov-Smirnov test shows that the ten H_2 -bearing systems with $\log f > -4.5$ (which is our conservative completeness limit) have H I column densities that are compatible with those of the overall DLA population. This may be due however to small number statistics. There is evidence (see Sect. 4) that the probability of finding large molecular fractions is higher in DLAs with large $N(\text{H I})$, as observed in the LMC. About 7% of the systems with $\log N(\text{H I}) < 20.8$ have $\log f > -4.5$ while $\sim 19\%$ of the systems with larger H I column densities have similar molecular fractions. There is no sharp transition in the molecular fraction of DLAs at any of the measured total hydrogen ($\text{H I} + \text{H}_2$) column density.

It is surprising to see that most of the DLAs have very low molecular fractions, i.e., $\log f \lesssim -6$. This is much smaller than observed along lines of sight in the Galactic disk (Savage et al. 1977). However, Wakker (2006) measured similar molecular fractions towards high-latitude Galactic lines of sight.

We confirm that a good criterion to find H_2 -bearing DLAs is to select high-metallicity systems. Indeed, 35% of the systems with $[\text{X}/\text{H}] \geq -1.3$ have $\log f > -4.5$ whilst this is the case for only 4% of those with metallicities lower than that. Since there is a correlation between metallicity and depletion factor, the latter being defined as $[\text{X}/\text{Fe}]$ with $\text{X} = \text{Zn}, \text{S}$ or Si , H_2 is found in DLAs also having the highest depletion factors. Therefore, clouds with large molecular fractions are expected to be dusty and clumpy (e.g., Hirashita et al. 2003). They can be missed because of the associated extinction and/or because of their small cross-section.

The presence of H_2 is closely related to the dust column density. Indeed, all detections pertain to systems where the column density of iron into dust grains is larger than $5 \times 10^{14} \text{ cm}^{-2}$, and about 40% of these systems have detectable H_2 . This shows that the presence of dust is an important ingredient in the formation of H_2 in DLAs. The low molecular fractions measured for most of the DLAs are probably a consequence of low abundances of metals and dust.

We show that the probability of finding H_2 increases with increasing velocity width of low-ionisation metal line profiles. The correlation between velocity width and metallicity (Ledoux et al. 2006a), if interpreted as a mass-metallicity relation, provides a natural explanation in which H_2 -bearing DLA systems are preferably associated with massive objects where star-formation is enhanced. There is no evidence of systematic outflows in H_2 -bearing DLAs neither from the H_2 profiles nor from the study of the associated high-ionisation phase.

From the comparison between H_2 -bearing systems and the overall UVES sample, we show that there is no evolution with redshift of the fraction of H_2 -bearing DLAs nor of the molecular fraction in systems with detected H_2 over the range $1.8 < z_{\text{abs}} \leq 3$. This, compared to the large amounts of H_2 observed in the local Universe, suggests that a significant increase of the molecular fraction in DLAs could take place at redshifts $z_{\text{abs}} \leq 1.8$. Ultraviolet observations from space are therefore needed to observe the H_2 Lyman and Werner bands in low and intermediate redshift DLAs.

Increasing the sample of H₂-bearing DLAs is required to cover a large range in column density and derive the H₂ frequency distribution, $f(N(\text{H}_2), X)$ – where X is the absorption distance – in a way similar to studies completed for H I (Lanzetta et al. 1991). This would assess whether the steep slope observed in $f(N(\text{H I}), X)$ at large column densities (e.g., Prochaska et al. 2005) is due to the conversion H I → H₂ (Schaye 2001) rather than to a magnitude-dependent bias from dust obscuration of the background quasars (e.g., Boissé et al. 1998; Vladilo & Péroux 2005; Smette et al. 2005). Both effects can explain the apparent lack of DLAs of both high metallicity and large H I column density, as observed in all DLA samples.

While the exact nature of DLAs is still an open debate, H₂-bearing DLAs provide interesting probes of the physical conditions close to star-forming regions. About a decade ago, only one H₂-bearing DLA was known. With the criteria revealed by the present survey, it will be possible to select these systems more efficiently and their numbers should increase rapidly. This opens up the exciting prospect of the detailed study of the ISM in distant galaxies.

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Table 1. Molecular and metal contents of UVES DLAs/sub-DLAs at $1.8 < z_{\text{abs}} \leq 4.2$.

Quasar B1950	z_{em}	z_{abs}	$\log N(\text{H I})$	$[\text{X}/\text{H}]$	$[\text{X}/\text{Fe}]$	X	ΔV [km s^{-1}]	$\log N(\text{H}_2)$		$\log f$	Refs.
								J = 0	J = 1		
Q 0000–2619	4.11	3.390	21.40 ± 0.08	-2.06 ± 0.09	-0.03 ± 0.06	Zn	33	< 13.9	≤ 13.8	< -6.87	<i>a, b</i>
Q 0010–0012	2.15	2.025	20.95 ± 0.10	-1.43 ± 0.11	-0.16 ± 0.06	Zn	32	< 14.4	< 14.5	< -5.80	<i>c</i>
Q 0013–0029	2.09	1.973	20.83 ± 0.05	-0.59 ± 0.05	0.83 ± 0.01	Zn	720	$18.86^{+1.14}_{-1.04}$	$17.30^{+0.07}_{-0.07}$	$-1.68^{+1.07}_{-0.18}$	<i>c, d, e</i>
Q 0027–1836	2.56	2.402	21.75 ± 0.10	-1.63 ± 0.10	0.65 ± 0.03	Zn	44	$17.30^{+0.07}_{-0.07}$		$-4.15^{+0.18}_{-0.17}$	<i>f, g</i>
Q 0039–3354	2.48	2.224	20.60 ± 0.10	-1.31 ± 0.12	0.38 ± 0.06	Si	122	< 13.8	< 13.9	< -6.05	<i>g, h</i>
Q 0049–2820	2.26	2.071	20.45 ± 0.10	-1.31 ± 0.12	0.14 ± 0.06	Si	51	< 13.8	< 14.0	< -5.84	<i>g, h</i>
Q 0058–2914	3.09	2.671	21.10 ± 0.10	-1.53 ± 0.10	0.32 ± 0.03	Zn	34	< 13.7	< 14.2	< -6.38	<i>c</i>
Q 0100+1300	2.69	2.309	21.35 ± 0.08	-1.58 ± 0.08	0.22 ± 0.01	Zn	37	< 13.5	< 13.9	< -6.92	<i>h</i>
Q 0102–1902	3.04	2.370	21.00 ± 0.08	-1.90 ± 0.08	0.16 ± 0.03	S	17	< 14.2	< 14.8	< -5.72	<i>c</i>
Q 0102–1902	3.04	2.926	20.00 ± 0.10	-1.50 ± 0.10	0.20 ± 0.03	Si	146	< 12.8	< 13.0	< -6.39	<i>h</i>
Q 0112–3030	2.99	2.418	20.50 ± 0.08	-2.42 ± 0.08	0.16 ± 0.04	Si	31	< 14.1	< 14.3	< -5.61	<i>c</i>
Q 0112–3030	2.99	2.702	20.30 ± 0.10	-0.49 ± 0.11	0.51 ± 0.08	Si	218	< 14.1	< 14.0	< -5.55	<i>c</i>
Q 0112+0259	2.81	2.423	20.90 ± 0.10	-1.31 ± 0.11	0.24 ± 0.04	S	112	< 14.1	< 14.2	< -6.05	<i>c</i>
Q 0131+0345	4.15	3.692	20.70 ± 0.10	< -2.55	< 0.08	S	39	< 13.7	< 13.9	< -6.19	<i>h, †</i>
Q 0131+0345	4.15	3.774	20.60 ± 0.10	-0.70 ± 0.10	0.19	Si	128	< 14.0	< 14.0	< -5.90	<i>h, *</i>
Q 0135–2722	3.21	2.800	21.00 ± 0.10	-1.40 ± 0.10	0.33 ± 0.04	S	65	< 14.3	< 14.1	< -6.09	<i>c</i>
Q 0216+0803	2.99	2.293	20.50 ± 0.10	-0.70 ± 0.11	0.42 ± 0.05	Zn	104	< 14.3	< 14.5	< -5.39	<i>i</i>
Q 0242–2917	3.27	2.560	20.90 ± 0.10	-1.99 ± 0.10	0.05 ± 0.04	S	70	< 14.5	< 14.2	< -5.82	<i>g, h</i>
Q 0254–4025	2.28	2.046	20.45 ± 0.08	-1.55 ± 0.09	0.23 ± 0.04	S	37	< 14.2	< 14.6	< -5.32	<i>g, h</i>
Q 0300–3152	2.37	2.179	20.80 ± 0.10	-1.80 ± 0.11	0.29 ± 0.04	S	41	< 14.5	< 14.4	< -5.65	<i>g, h</i>
Q 0331–4506	2.67	2.411	20.15 ± 0.07	< -1.41	< 0.08	Si	31	< 14.6	< 14.8	< -4.77	<i>g, h</i>
Q 0335–1213	3.44	3.180	20.65 ± 0.10	-2.44 ± 0.14	0.01 ± 0.16	Si	15	< 14.9	< 14.8	< -5.10	<i>h, j</i>
Q 0336–0143	3.20	3.062	21.10 ± 0.10	-1.41 ± 0.10	0.40 ± 0.03	Si	67	< 13.5	< 13.5	< -6.90	<i>h</i>
Q 0347–3819	3.22	3.025	20.73 ± 0.05	-1.17 ± 0.07	0.71 ± 0.05	Zn	93	$14.53^{+0.06}_{-0.06}$		$-5.90^{+0.11}_{-0.11}$	<i>c, k</i>
Q 0405–4418	3.02	2.550	21.15 ± 0.15	-1.36 ± 0.16	0.24 ± 0.04	Zn	165	< 13.9	< 13.6	< -6.62	<i>c</i>
Q 0405–4418	3.02	2.595	21.05 ± 0.10	-1.12 ± 0.10	0.34 ± 0.05	Zn	79	$18.14^{+0.07}_{-0.10}$		$-2.61^{+0.17}_{-0.20}$	<i>c, l</i>
Q 0405–4418	3.02	2.622	20.45 ± 0.10	-2.04 ± 0.10	0.32 ± 0.02	Si	182	< 13.3	< 13.7	< -6.20	<i>c</i>
Q 0421–2624	2.28	2.157	20.65 ± 0.10	-1.86 ± 0.10	0.32 ± 0.01	Si	47	< 13.7	< 13.9	< -6.14	<i>g, h</i>
Q 0425–5214	2.25	2.224	20.30 ± 0.10	-1.43 ± 0.11	0.41 ± 0.04	S	40	< 13.9	< 14.1	< -5.59	<i>g, h</i>
Q 0432–4401	2.65	2.302	20.95 ± 0.10	-1.23 ± 0.13	0.35 ± 0.14	Si	88	< 15.1	< 15.1	< -5.15	<i>h, j</i>
Q 0450–1310	2.25	2.067	20.50 ± 0.07	-1.62 ± 0.08	0.14 ± 0.04	S	148	< 13.5	< 13.9	< -6.08	<i>h</i>
Q 0458–0203	2.29	2.040	21.70 ± 0.10	-1.22 ± 0.10	0.65 ± 0.04	Zn	88	< 14.6	< 14.6	< -6.40	<i>i, m</i>
Q 0528–2505	2.77	2.811	21.35 ± 0.07	-0.91 ± 0.07	0.46 ± 0.01	Zn	304	$18.22^{+0.11}_{-0.12}$		$-2.83^{+0.18}_{-0.19}$	<i>b, l, n, o, p</i>
Q 0551–3638	2.32	1.962	20.70 ± 0.08	-0.35 ± 0.08	0.80 ± 0.04	Zn	468	$17.42^{+0.45}_{-0.73}$		$-2.98^{+0.33}_{-0.81}$	<i>c, q</i>
Q 0642–5038	3.09	2.659	20.95 ± 0.08	-1.05 ± 0.09	0.30 ± 0.04	Zn	99	$18.41^{+0.03}_{-0.04}$		$-2.24^{+0.11}_{-0.12}$	<i>h</i>
Q 0841+1256	2.50	2.375	21.05 ± 0.10	-1.59 ± 0.10	0.24 ± 0.03	Zn	37	≤ 14.56	< 14.0	< -5.98	<i>b</i>
Q 0841+1256	2.50	2.476	20.80 ± 0.10	-1.60 ± 0.10	0.22 ± 0.04	Zn	30	< 13.7	< 13.9	< -6.29	<i>b</i>
Q 0913+0714	2.78	2.618	20.35 ± 0.10	-2.59 ± 0.10	0.16 ± 0.03	Si	22	< 13.7	< 13.8	< -5.90	<i>h</i>
Q 0933–3319	2.91	2.682	20.50 ± 0.10	-1.27 ± 0.14	0.28 ± 0.14	Si	312	< 14.9	< 14.4	< -5.08	<i>h, j</i>
Q 0951–0450	4.37	4.203	20.55 ± 0.10	-2.68 ± 0.10	> 0.04	Si	15	< 13.2	< 13.3	< -6.60	<i>h, †</i>
Q 1036–2257	3.13	2.778	20.93 ± 0.05	-1.36 ± 0.05	0.31 ± 0.01	S	80	< 13.0	< 13.0	< -7.28	<i>h</i>
Q 1108–0747	3.92	3.608	20.37 ± 0.07	-1.59 ± 0.07	0.40 ± 0.02	Si	31	< 14.0	< 14.2	< -5.59	<i>h</i>
Q 1111–1517	3.37	3.266	21.30 ± 0.05	-1.65 ± 0.11	0.34 ± 0.09	Zn	140	< 13.1	< 13.4	< -7.37	<i>h</i>
Q 1117–1329	3.96	3.351	20.95 ± 0.10	-1.41 ± 0.11	0.23 ± 0.06	Zn	43	< 14.0	< 13.6	< -6.40	<i>c</i>
Q 1157+0128	1.99	1.944	21.80 ± 0.10	-1.44 ± 0.10	0.39 ± 0.01	Zn	89	< 14.4	< 14.5	< -6.65	<i>b</i>
Q 1209+0919	3.30	2.584	21.40 ± 0.10	-1.01 ± 0.10	0.53 ± 0.03	Zn	214	< 14.9	< 15.1	< -5.69	<i>i</i>
Q 1220–1800	2.16	2.113	20.12 ± 0.07	-0.93 ± 0.07	0.34 ± 0.02	S	95	< 13.8	< 14.0	< -5.54	<i>g, h</i>
Q 1223+1753	2.94	2.466	21.40 ± 0.10	-1.63 ± 0.10	0.06 ± 0.02	Zn	91	< 13.9	< 14.1	< -6.69	<i>b</i>
Q 1232+0815	2.57	2.338	20.90 ± 0.08	-1.43 ± 0.08	0.45 ± 0.01	S	85	$19.57^{+0.10}_{-0.10}$		$-1.07^{+0.16}_{-0.17}$	<i>b, l, r, s</i>
Q 1337+1121	2.92	2.508	20.12 ± 0.05	-1.81 ± 0.06	0.41 ± 0.04	Si	32	< 13.8	< 14.2	< -5.42	<i>c</i>
Q 1337+1121	2.92	2.796	21.00 ± 0.08	-1.86 ± 0.09	0.31 ± 0.04	Si	42	< 13.7	< 13.5	< -6.71	<i>c</i>
Q 1340–1340	3.20	3.118	20.05 ± 0.08	-1.42 ± 0.08	0.20 ± 0.03	S	153	< 12.5	< 12.9	< -6.62	<i>h</i>
Q 1354–1046	3.01	2.501	20.44 ± 0.05	-1.51 ± 0.11	0.08 ± 0.14	S	71	< 14.7	< 14.5	< -5.18	<i>h, j</i>
Q 1354–1046	3.01	2.967	20.80 ± 0.10	-1.39 ± 0.10	0.28 ± 0.06	Si	30	< 13.1	< 13.5	< -6.75	<i>h, j</i>
Q 1418–0630	3.69	3.448	20.50 ± 0.10	-1.43 ± 0.13	0.39 ± 0.10	Si	35	< 14.0	< 14.2	< -5.69	<i>h, j</i>
Q 1441+2737	4.42	4.224	20.95 ± 0.10	-0.63 ± 0.10	0.49 ± 0.03	Zn	130	$18.29^{+0.07}_{-0.08}$		$-2.36^{+0.17}_{-0.18}$	<i>i, t</i>
Q 1444+0126	2.21	2.087	20.25 ± 0.07	-0.80 ± 0.09	0.95 ± 0.07	Zn	294	$18.16^{+0.14}_{-0.12}$		$-1.80^{+0.21}_{-0.19}$	<i>c</i>
Q 1451+1223	3.25	2.469	20.40 ± 0.10	-2.27 ± 0.13	0.25 ± 0.08	Si	35	< 15.2	< 15.2	< -4.50	<i>b</i>
Q 1451+1223	3.25	3.171	20.20 ± 0.20	-2.10 ± 0.21	0.25 ± 0.10	Si	45	< 13.6	< 13.6	< -5.80	<i>b</i>
Q 2059–3604	3.09	2.507	20.29 ± 0.07	-1.85 ± 0.20	0.44 ± 0.19	S	25	< 14.5	< 14.6	< -5.07	<i>c</i>
Q 2059–3604	3.09	3.083	20.98 ± 0.08	-1.77 ± 0.09	0.23 ± 0.04	S	44	< 13.4	< 13.5	< -6.85	<i>b</i>
Q 2116–3550	2.34	1.996	20.10 ± 0.07	-0.34 ± 0.11	0.54 ± 0.09	Zn	177	< 14.5	< 14.8	< -4.75	<i>i</i>
Q 2138–4427	3.17	2.852	20.98 ± 0.05	-1.74 ± 0.05	0.11 ± 0.03	Zn	58	< 13.4	< 13.6	< -6.82	<i>c</i>
Q 2206–1958	2.56	1.921	20.67 ± 0.05	-0.54 ± 0.05	0.29 ± 0.01	Zn	136	< 14.4	< 14.7	< -5.44	<i>i</i>
Q 2206–1958	2.56	2.076	20.44 ± 0.05	-2.32 ± 0.05	0.28 ± 0.02	Si	20	< 13.2	< 13.6	< -6.34	<i>h</i>
Q 2222–3939	2.18	2.154	20.85 ± 0.10	-1.97 ± 0.10	-0.04 ± 0.04	S	21	< 13.8	< 13.8	< -6.35	<i>g, h</i>
Q 2228–3954	2.21	2.095	21.20 ± 0.10	-1.36 ± 0.12	0.17 ± 0.07	Zn	138	< 13.8	< 14.0	< -6.59	<i>g, h</i>
Q 2230+0232	2.15	1.864	20.90 ± 0.10	-0.81 ± 0.10	0.24 ± 0.01	S	148	< 15.4	< 15.4	< -4.80	<i>i</i>
Q 2243–6031	3.01	2.331	20.65 ± 0.05	-0.85 ± 0.05	0.35 ± 0.02	Zn	173	< 13.8	$<$		

Table 1. continued.

Quasar B1950	z_{em}	z_{abs}	$\log N(\text{H I})$	$[\text{X}/\text{H}]$	$[\text{X}/\text{Fe}]$	X	ΔV [km s $^{-1}$]	$\log N(\text{H}_2)$		$\log f$	Refs.
								J = 0	J = 1		
Q 2343+1232	2.51	2.431	20.40 ± 0.07	-0.87 ± 0.10	0.51 ± 0.07	Zn	289	$13.69^{+0.09}_{-0.09}$	13.4	$-6.41^{+0.16}_{-0.16}$	<i>f, i</i>
Q 2344+1229	2.76	2.538	20.50 ± 0.10	-1.81 ± 0.10	0.18 ± 0.01	Si	69	< 13.4	< 13.4	< -6.40	<i>h</i>
Q 2348-0108	3.01	2.426	20.50 ± 0.10	-0.62 ± 0.10	0.55 ± 0.02	S	248	$18.52^{+0.29}_{-0.49}$	13.6	$-1.69^{+0.38}_{-0.58}$	<i>i, u</i>
Q 2348-0108	3.01	2.615	21.30 ± 0.08	-2.02 ± 0.08	0.13 ± 0.01	Si	95	< 13.1	< 13.6	< -7.20	<i>h, u</i>
Q 2348-1444	2.94	2.279	20.63 ± 0.05	-2.10 ± 0.08	0.25 ± 0.10	S	55	< 13.6	< 13.6	< -6.38	<i>h</i>

^a Levshakov et al. (2000, 2001), ^b Petitjean et al. (2000), ^c Ledoux et al. (2003), ^d Ge & Bechtold (1997), ^e Petitjean et al. (2002), ^f Noterdaeme et al. (2007a), ^g Smette et al. (2005), ^h this work, ⁱ Petitjean et al. (2006), ^j Akerman et al. (2005), ^k Levshakov et al. (2002), ^l Srianand et al. (2005), ^m Heinmüller et al. (2006), ⁿ Levshakov & Varshalovich (1985), ^o Foltz et al. (1988), ^p Srianand & Petitjean (1998), ^q Ledoux et al. (2002), ^r Srianand et al. (2000), ^s Ge et al. (2001), ^t Ledoux et al. (2006b), ^u Noterdaeme et al. (2007b).

[†] $N(\text{Fe II})$ from Prochaska et al. (2007b).

* Metal column densities from Prochaska et al. (2007b). Only an upper limit on $N(\text{Fe II})$ could be measured (Prochaska et al. 2003). The abundance of iron is therefore estimated through that of nickel, assuming $[\text{Fe}/\text{H}] = [\text{Ni}/\text{H}] - 0.1$ dex.